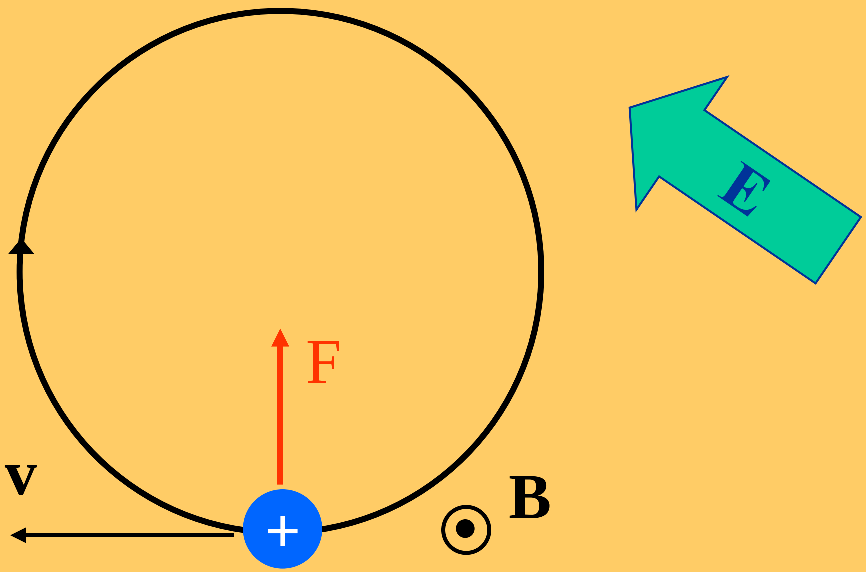
Think about this:

$$\mathbf{F} = q\mathbf{v} \times \mathbf{B}$$



What happens if you add an electric field **E**?

Last lecture (3)

- Solar activity

Today's lecture (4)

- Solar wind – basic facts
- Solar wind – magnetic structure
- Ionosphere

  - layers

  - radio wave reflection

Today

Activity	Date	Time	Room	Subject	Litterature
L1	2/9	10-12	Q33	Course description, Introduction, The Sun 1, Plasma physics 1	**CGF** Ch 1, 5, (p 110-113)
L2	4/9	10-12	Q21	The Sun 2, Plasma physics 2	**CGF** Ch 5 (p 114-121), 6.3
L3	8/9	13-15	Q36	Solar wind, The ionosphere and atmosphere 1, Plasma physics 3	**CGF** Ch 6.1, 2.1-2.6, 3.1-3.2, 3.5, **LL** Ch III, Extra material
T1	10/9	10-12	Q33	Mini-group work 1	
L4	15/9	13-15	Q31	The ionosphere 2, Plasma physics 4	**CGF** Ch 3.4, 3.7, 3.8
T2	17/9	10-12	Q33	Mini-group work 2	
L5	19/9	15-17	Q31	The Earth's magnetosphere 1, Plasma physics 5	**CGF** 4.1-4.3, **LL** Ch I, II, IV.A
L6	23/9	8-10	Q31	The Earth's magnetosphere 2, Other magnetospheres	**CGF** Ch 4.6-4.9, **LL** Ch V.
T3	24/9	14-16	Q21	Mini-group work 3	
L7	29/9	11-13	Q36	Aurora, Measurement methods in space plasmas and data analysis 1	**CGF** Ch 4.5, 10**, LL** Ch VI, Extra material
T4	1/10	15-17	Q31	Mini-group work 4	
L8	2/10	15-17	Q34	Space weather and geomagnetic storms	**CGF** Ch 4.4, **LL** Ch IV.B-C, VII.A-C
L9	8/10	13-15	Q36	Interstellar and intergalactic plasma, Cosmic radiation, Swedish and international space physics research.	**CGF** Ch 7-9
T5	9/10	15-17	Q31	Mini-group work 5	
L10	13/10	15-17	Q33	Guest lecture (preliminary): Swedish astronaut Christer Fuglesang	
T6	16/10	10-12	Q36	Round-up	
Written exami-nation	30/10	8-13	M33, M37, M38		

Rosetta mission

Comet 67P/Churyumov-Gerasimenko

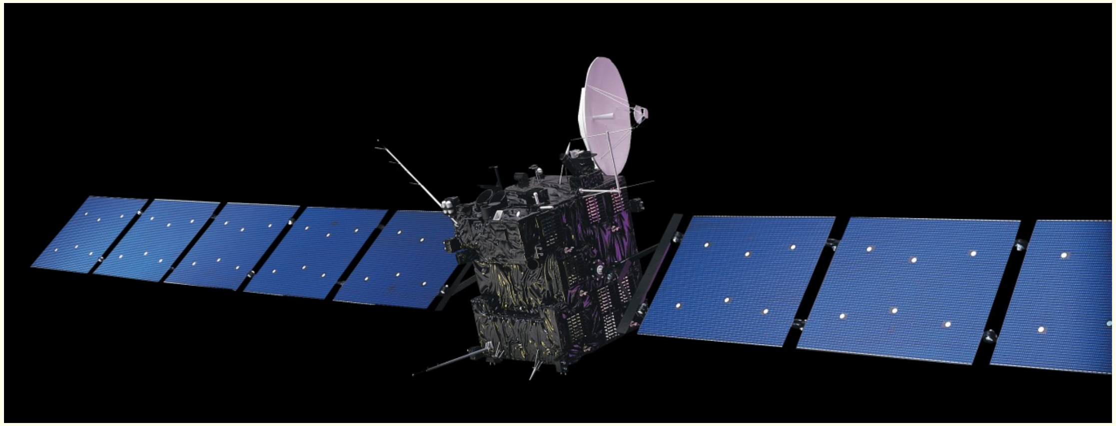
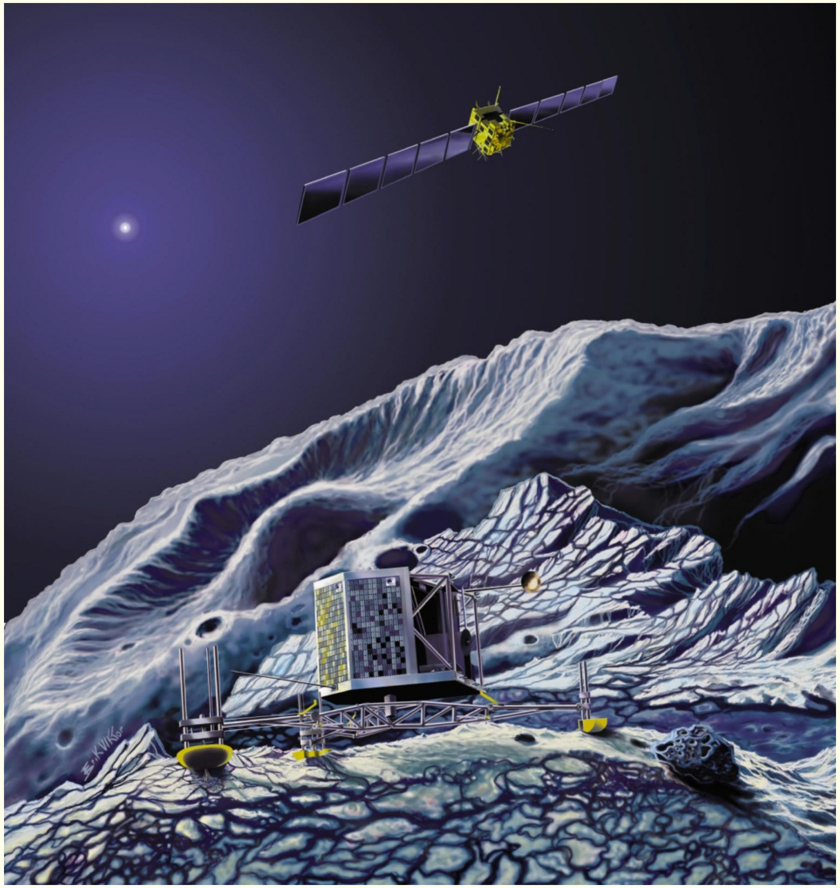


Photo from the Giotto spacecraft, ca 1985.

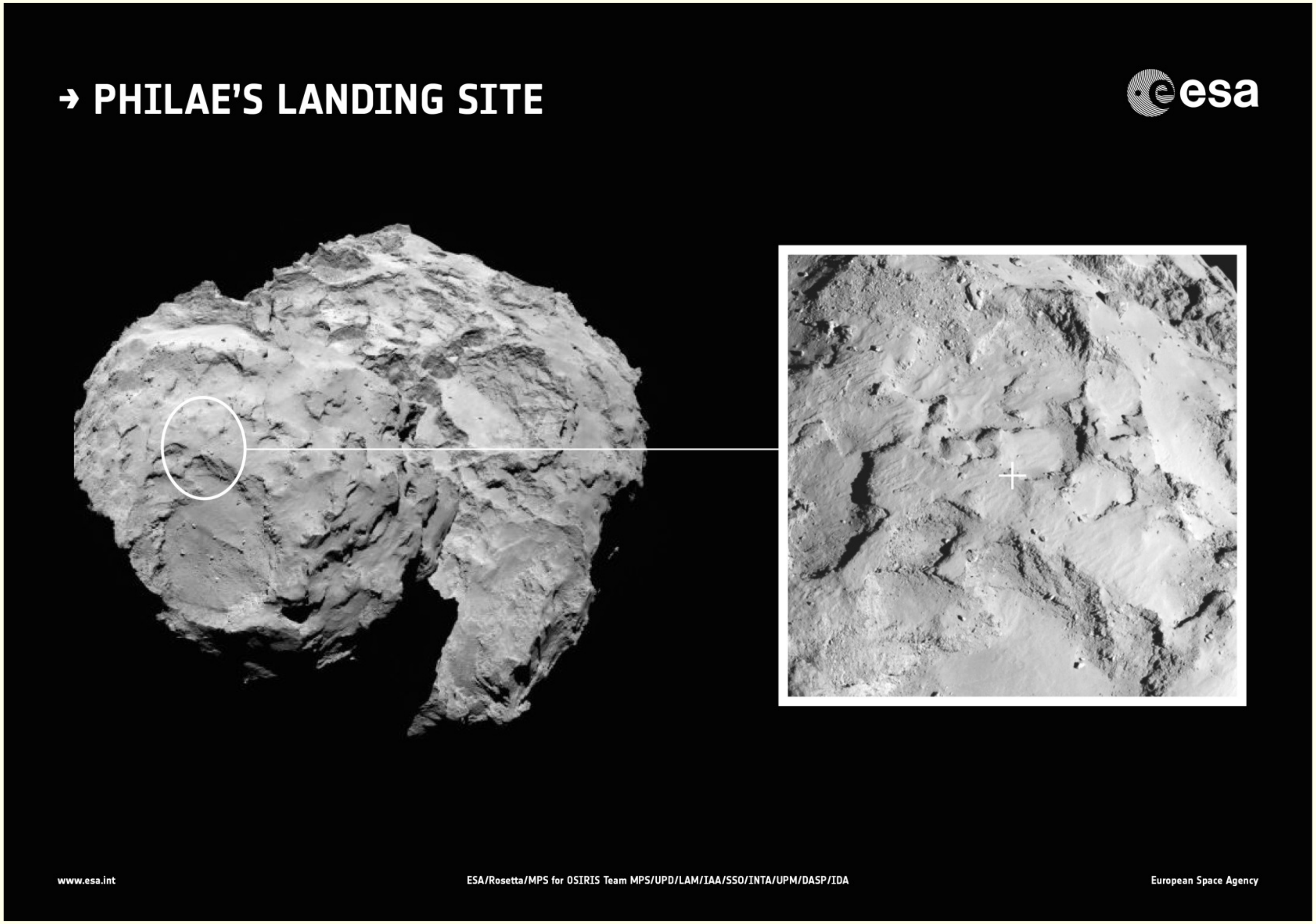


Launched, Feb., 2004, arrival August 2014.
KTH involved in plasma density instrument.

Landing site announced today!

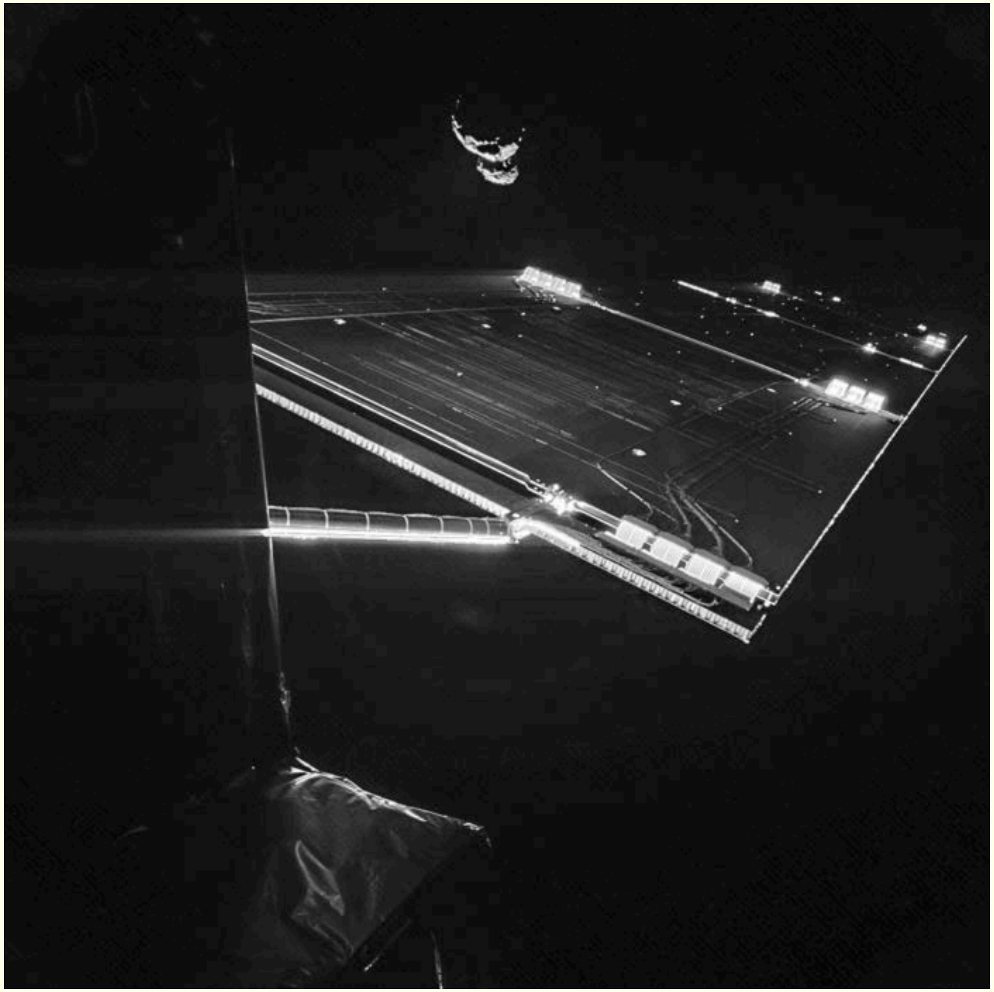
Rosetta mission

Philea landing site decided today!



Rosetta mission

Comet 67P/Churyumov-Gerasimenko



Mini groupwork 1

a)

$$h = \frac{42\ \text{mm}}{7\ \text{mm}} \cdot 6378\ \text{km} \cdot 2 = 77000\ \text{km}$$

The thermal energy is divided into motion in the three dimensions, two of which only give rise to a gyro motion around the magnetic field lines, with the motion along the magnetic field corresponding to an energy

$$E = \frac{k_B T}{2} = \frac{1.38 \cdot 10^{-23} \cdot 1.5 \cdot 10^{6}}{2} = 1 \cdot 10^{-17}\ \text{J}$$

$$v = \sqrt{\frac{2E}{m_e}} = \sqrt{\frac{2 \cdot 10^{-17}}{0.91 \cdot 10^{-30}}} = 4.7 \cdot 10^{6}\ \text{ms}^{-1}$$

Approximating the loop with a quarter-circle, the electron has to travel a length

$s = \pi h/2$ = 120 000 km

Then we get t = 25 s.

Energy - temperature

Average energy of molecule/atom:

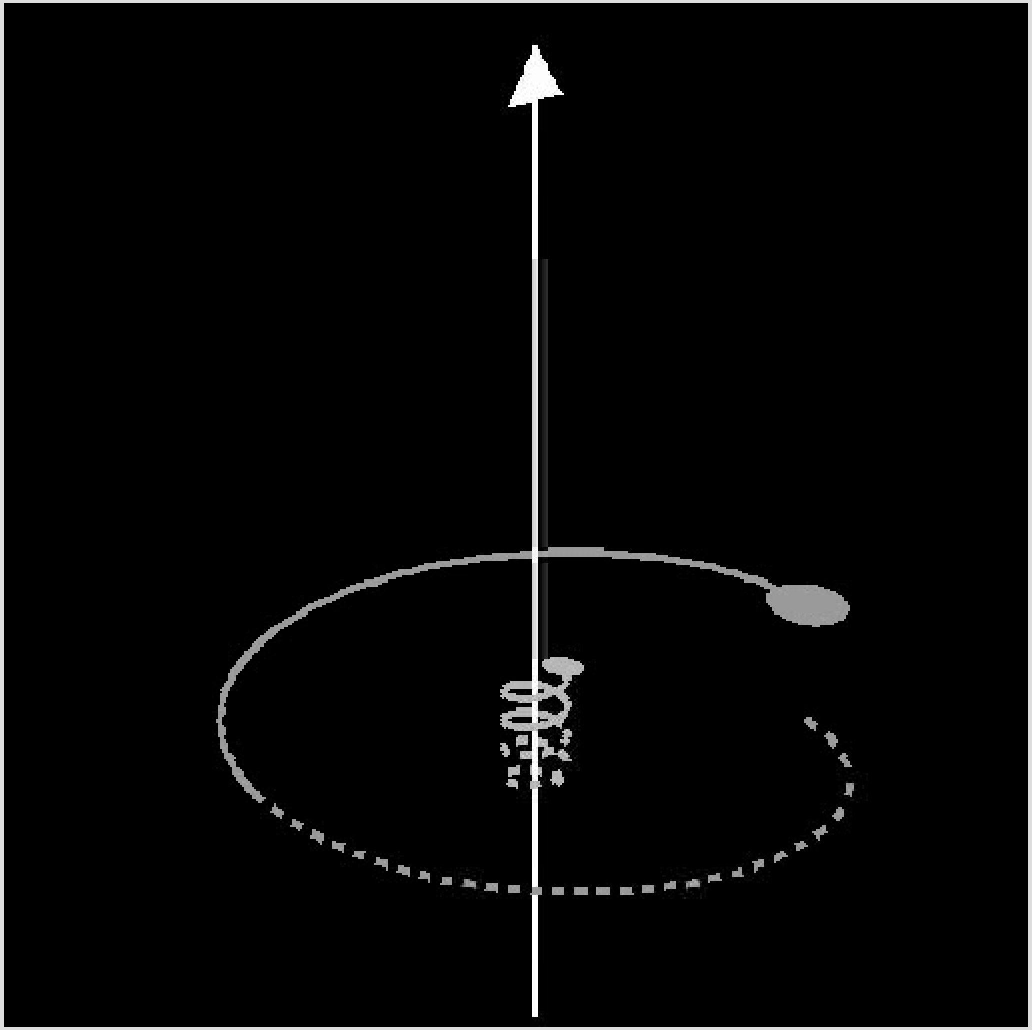
$$E = \frac{3}{2} k_B T \quad \Rightarrow$$

$$T = \frac{2E}{3k_B}$$

1 eV = 1.6·10^{-19} J ⇒

$$T = \frac{2E}{3k_B} = \frac{2 \cdot 1.6 \cdot 10^{-19}\,\mathrm{J}}{3 \cdot 1.38 \cdot 10^{-23}\,\dfrac{\mathrm{J}}{\mathrm{K}}} = 7729\,\mathrm{K}$$

Gyro motion



Equipartion principle

Statistically the kinetic energy is equally distributed along the three dimensions:

$$E_{\parallel} = \frac{1}{2} k_B T$$

$$E_{\perp} = \frac{2}{2} k_B T$$

Mini groupwork 1

b)

$$f_c = \frac{\omega_c}{2\pi} = \frac{1}{2\pi}\frac{qB}{m} \Rightarrow$$

$$B = \frac{2\pi f_c m}{q} = \frac{2\pi \cdot 1\cdot 10^{10} \cdot 0.91\cdot 10^{-30}}{1.6\cdot 10^{-19}} = 0.36\,\mathrm{T}$$

The perpendicular energy is given by

$$E = 2\cdot\frac{k_B T}{2} = 2\cdot\frac{1.38\cdot 10^{-23}\cdot 1.5\cdot 10^{6}}{2} = 2\cdot 10^{-17}\,\mathrm{J}$$

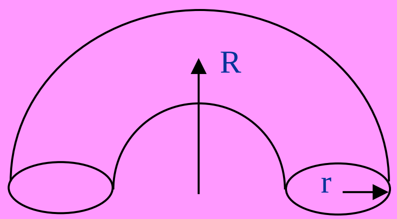
$$\Rightarrow$$

$$v = \sqrt{\frac{2E}{m_e}} = \sqrt{\frac{4\cdot 10^{-17}}{0.91\cdot 10^{-30}}} = 6.6\cdot 10^{6}\ \mathrm{ms^{-1}}$$

$$\rho = \frac{m_e v_\perp}{qB} = \frac{0.91\cdot 10^{-30}\cdot 6.6\cdot 10^{6}}{1.6\cdot 10^{-19}\cdot 0.36} = 1.0\cdot 10^{-4}\ \mathrm{m}$$

Mini groupwork 1

c) Model the flare by a half torus with minor axis r, and major axis. From the figure, estimate R = 2.6 R_E, and r = 2 R_E.

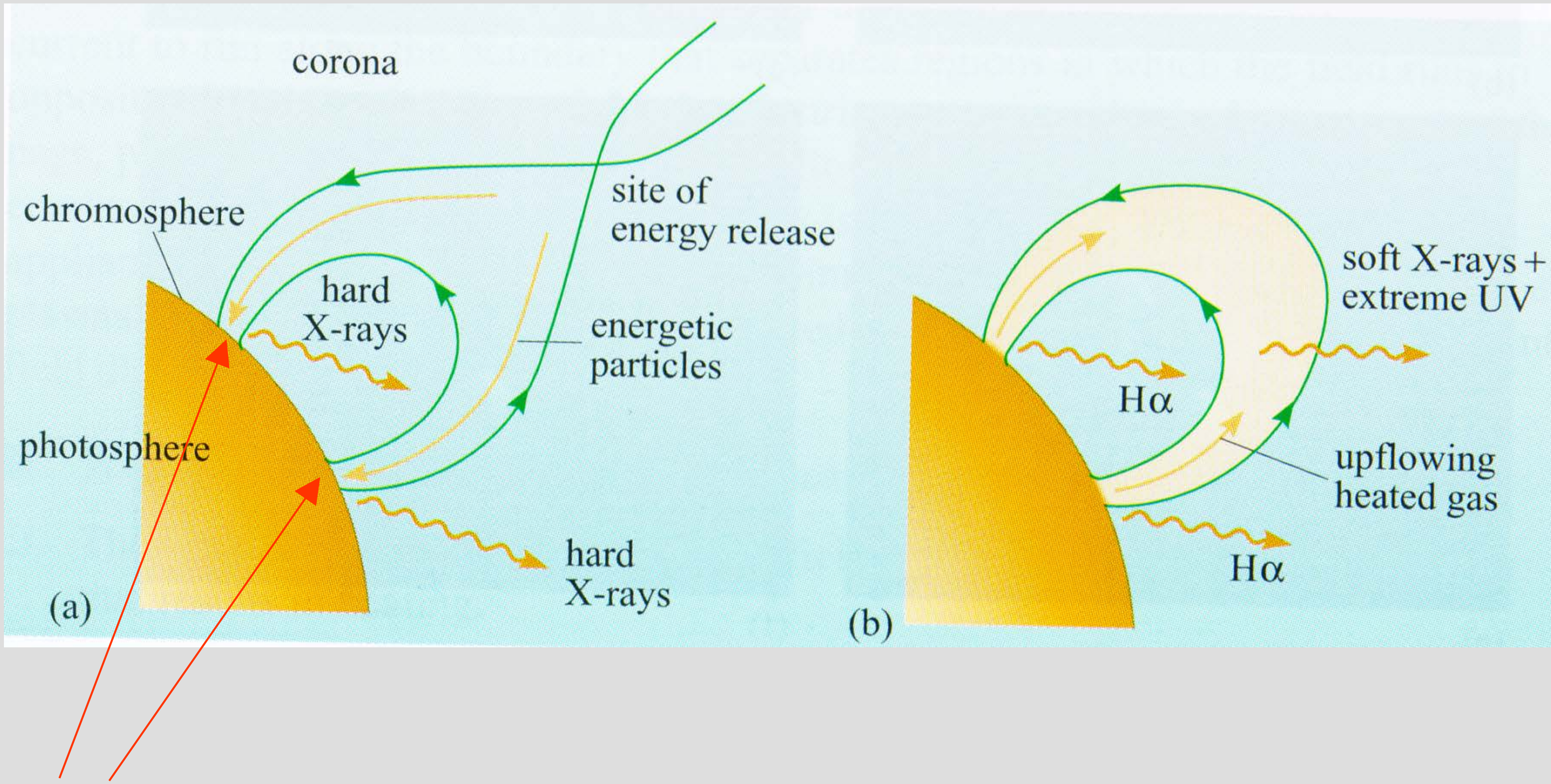


Let this half-torus be filled with a magnetic field of strength B ~ 0.36 T (using the value in b)). If the volume of the half-torus is V and the magnetic energy density is p_B, the total energy is

$$W = Vp_B = \pi R \pi r^2 \frac{B^2}{2\mu_0} = \pi^2 \cdot 2 \cdot 12^2 {R_E}^3 \frac{B^2}{2\mu_0}$$

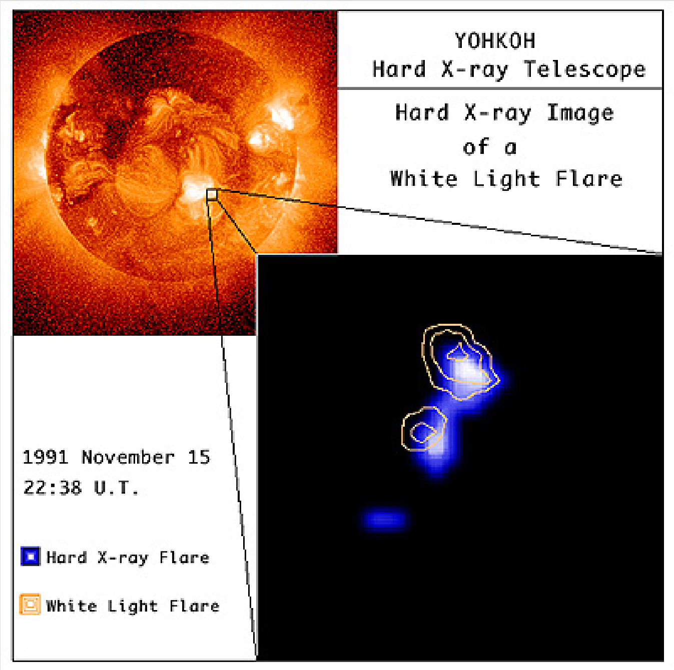
$$= \pi^2 \cdot 2 \cdot 12^2 \, (6378 \cdot 10^3)^3 \frac{(0.36)^2}{2\mu_0} = 3.8 \cdot 10^{28} \text{ J}$$

Solar flare mechanism

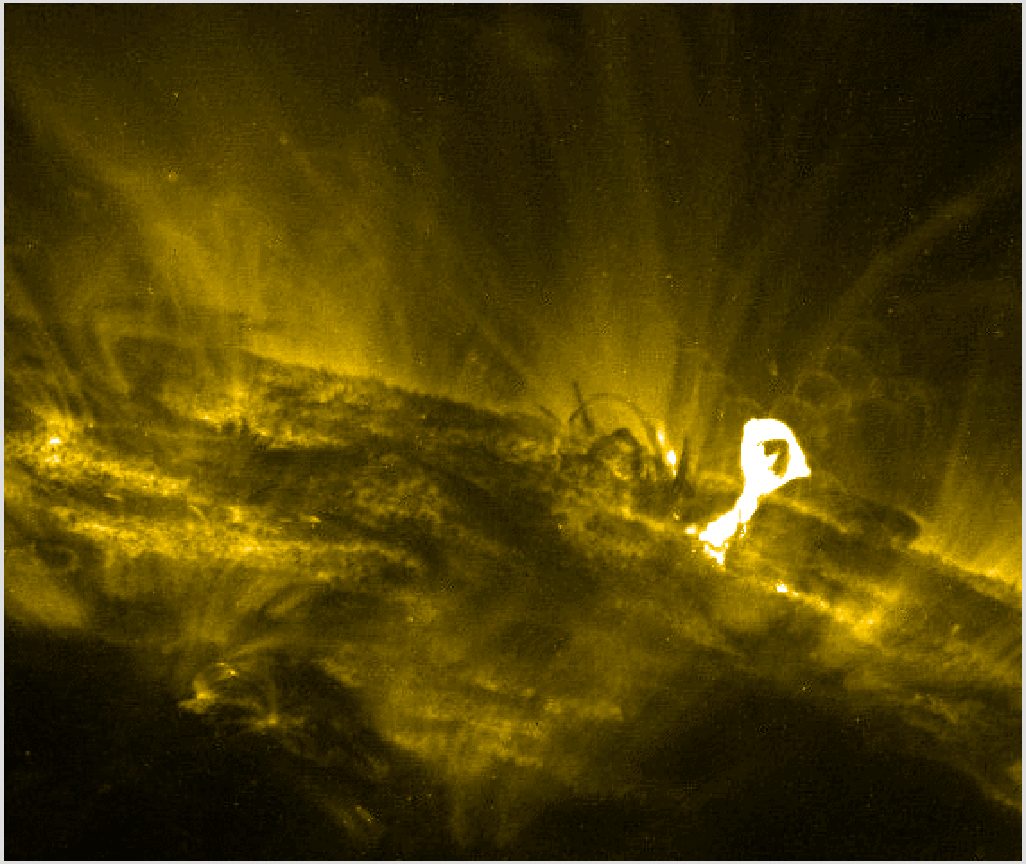


Electrons are accelerated, collide with solar surface (photosphere) and emit bremsstrahlung (X-rays).

Solar flare observations



(a) double signature of x-ray emissions at foot of flare



(b) coronal loop filled with hot gas

Frozen in magnetic flux *PROOF II*

$$\frac{\partial \mathbf{B}}{\partial t} = \underbrace{\nabla \times (\mathbf{v} \times \mathbf{B})}_{A} + \underbrace{\frac{1}{\mu_0 \sigma} \nabla^2 \mathbf{B}}_{B}$$

Order of magnitude estimate:

$$\frac{A}{B} = \frac{\nabla \times (\mathbf{v} \times \mathbf{B})}{\frac{1}{\mu_0 \sigma} \nabla^2 \mathbf{B}} \approx \frac{\frac{v \Delta B}{L}}{\frac{\Delta B}{\mu_0 \sigma L^2}} = v L \mu_0 \sigma \equiv R_m$$

Magnetic Reynolds number R_m:

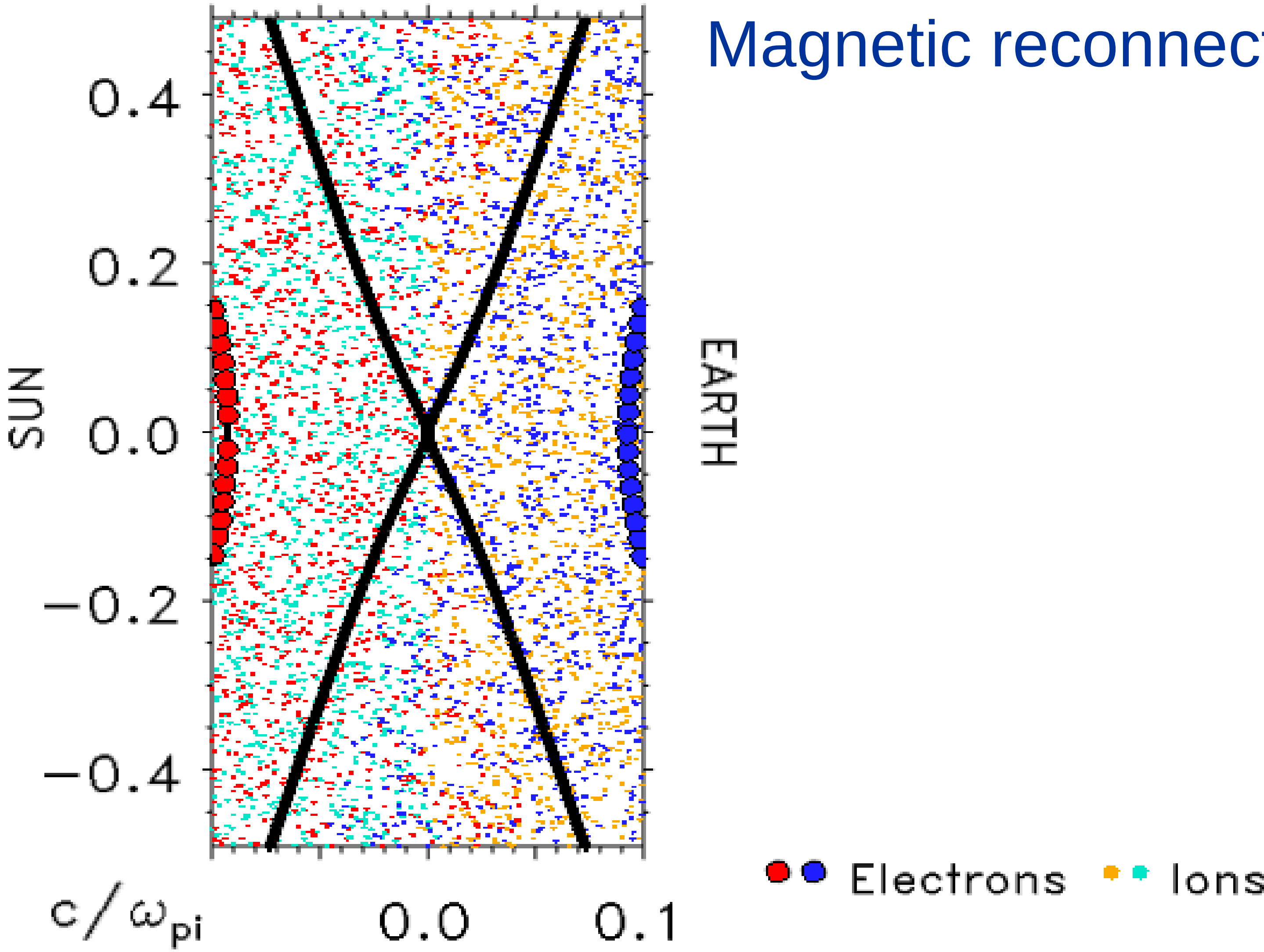
$$\mathrm{R_m} >> 1 \Rightarrow \quad \frac{\partial \mathbf{B}}{\partial t} = \nabla \times (\mathbf{v} \times \mathbf{B})$$

Frozen-in fields!

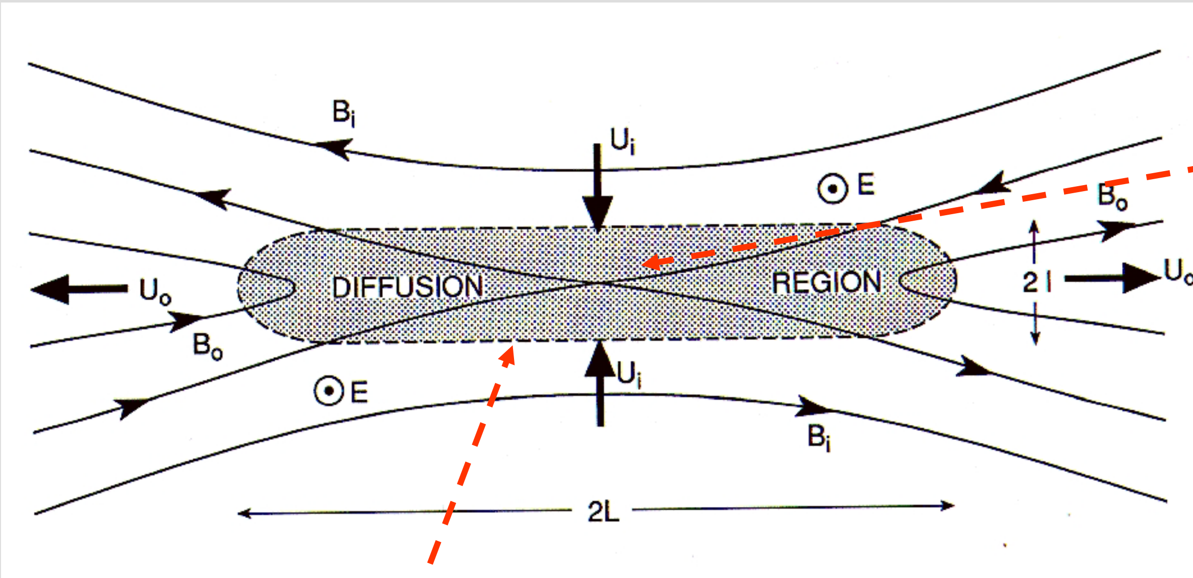
$$\mathrm{R_m} << 1 \Rightarrow \quad \frac{\partial \mathbf{B}}{\partial t} = \frac{1}{\mu_0 \sigma} \nabla^2 \mathbf{B}$$

Diffusion equation!

Magnetic reconnection



Reconnection



In 'diffusion region':

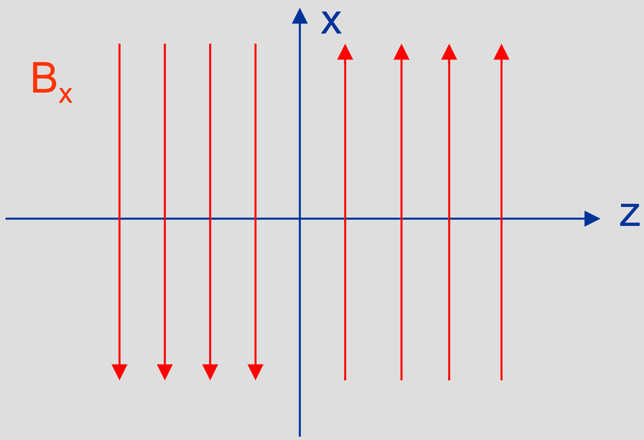
$R_m = \mu_0 \sigma l v \sim 1$

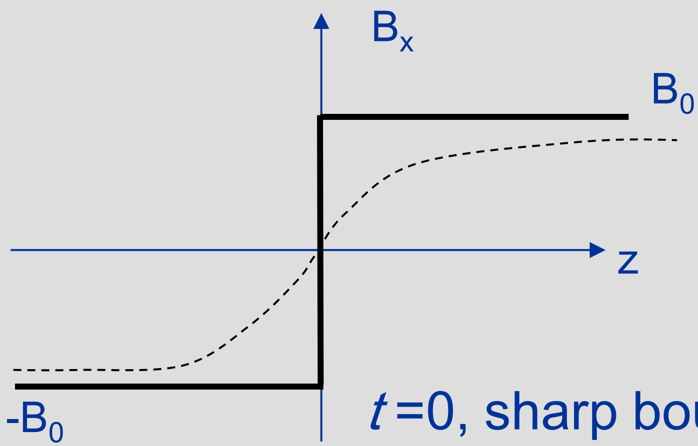
Thus: condition for frozen-in magnetic field breaks down.

A second condition is that there are two regions of magnetic field pointing in *opposite* direction:

- Field lines are "cut" and can be re-connected to other field lines
- **Magnetic energy is transformed into kinetic energy ($U_o >> U_i$)**
- **Plasma from different field lines can mix**

Reconnection in 1D





t =0, sharp boundary

$$\frac{\partial \mathbf{B}}{\partial t} = \frac{1}{\mu_0 \sigma} \nabla^2 \mathbf{B} \longrightarrow \frac{\partial B_x}{\partial t} = \frac{1}{\mu_0 \sigma} \frac{\partial^2 B_x}{\partial z^2}$$

Diffusion equation! Has solution

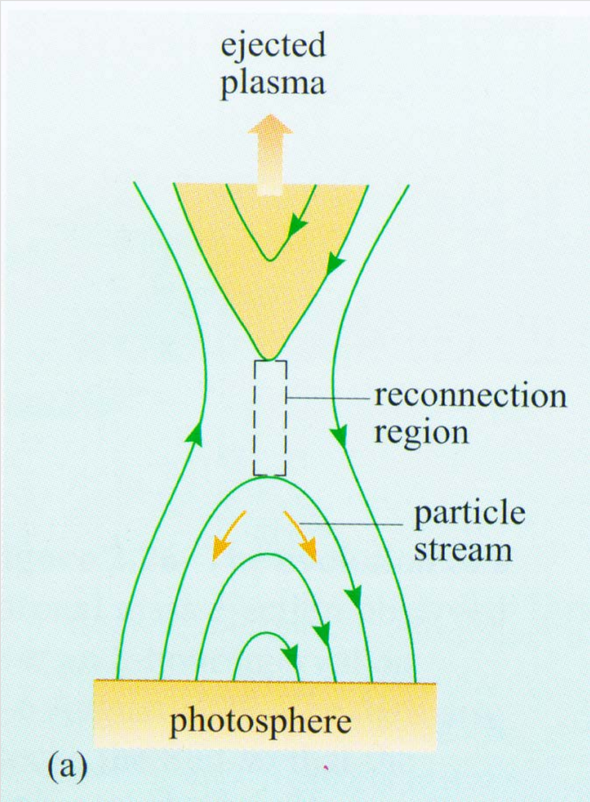
$$B_x(z,t) = B_0 erf\left(\left[\frac{\mu_0 \sigma}{4t}\right]^{1/2} z\right)$$

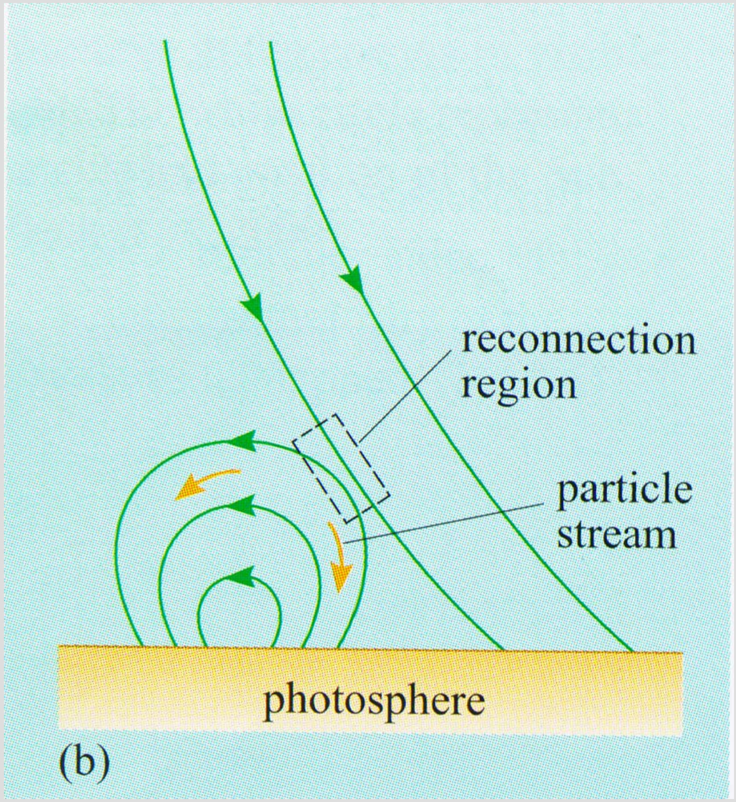
The total magnetic energy then decreases with time:

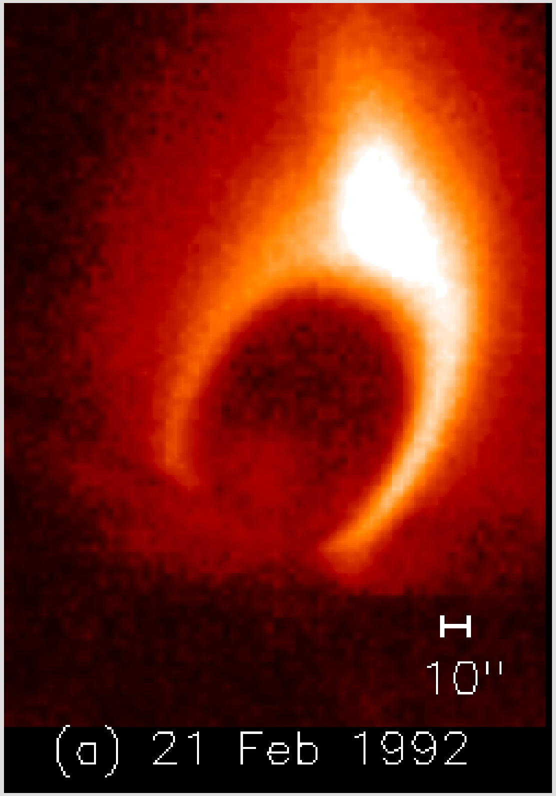
$$W_B = \int_{-\infty}^{\infty} \frac{B^2}{2\mu_0} dz$$

The magnetic energy is converted into heat and kinetic energy in 2D

Solar flare
energization mechanism

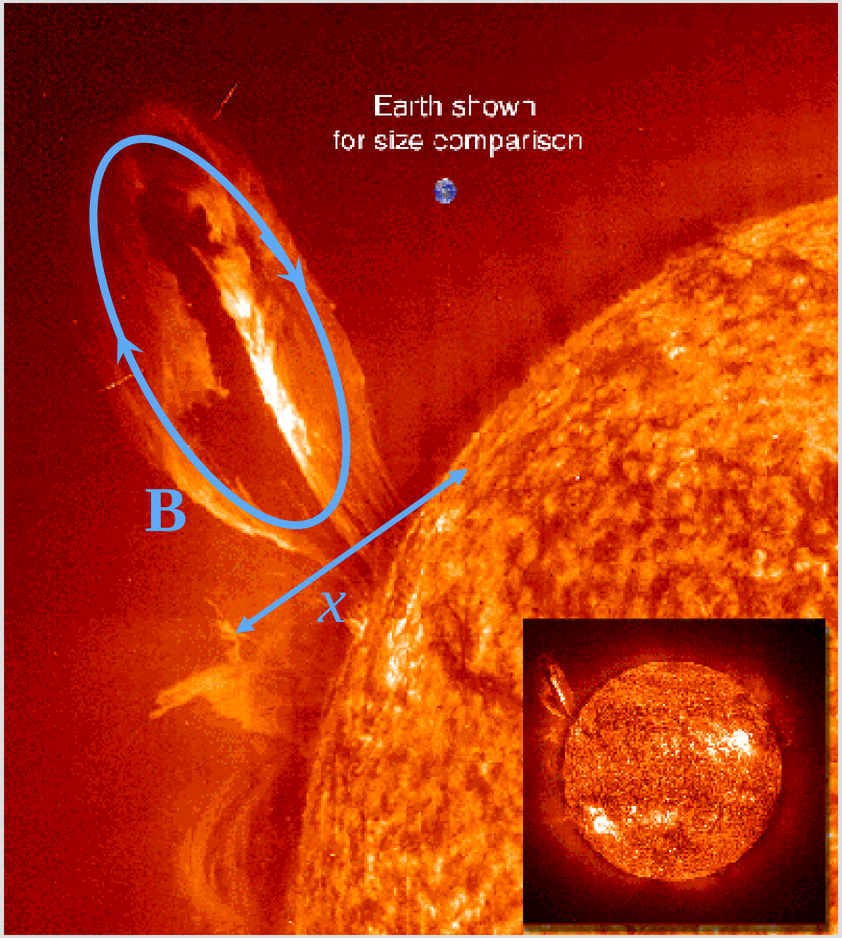






Two possible reconnection geometries

Coronal mass ejections



CME are sometimes called "magnetic clouds", because of their magnetic field configuration.

NIKLAS WAHLLÖF
"Så gott som allt upplevt – en frukost, ett bad och givetvis alla resor – är magiskt. Helt magiskt."

Läs Dagens Nyheter för endast 99 kr! Beställ här! NY APP!

Solstorm på väg mot jorden

Publicerad i dag 14:04

Foto: TT I solens atmosfär sker hela tiden olika fenomen och explosioner. Då och då sker urladdningar av magnetisk energi vilket skapar flammor av ultraviolett ljus och moln som består av elektriskt ledande gas.

En solstorm är just nu på väg mot jorden och kan träffa atmosfären i kväll eller i morgon. Det kan innebära norrsken över stora delar av Sverige.

I solens atmosfär sker hela tiden olika fenomen och explosioner. Då och då sker urladdningar av magnetisk energi vilket skapar flammor av ultraviolett ljus och kan kasta ut moln som består av elektriskt ledande gas.

I folkmun kallas det för solstormar och just nu är en sådan på väg rakt mot jorden. Nyhetsbyrån AP skriver att det var det flera år sedan en solstorm i den här storleken senast riskerade att träffa jorden, men enligt astronomen Dan Kiselman innebär det ingen fara för oss människor.

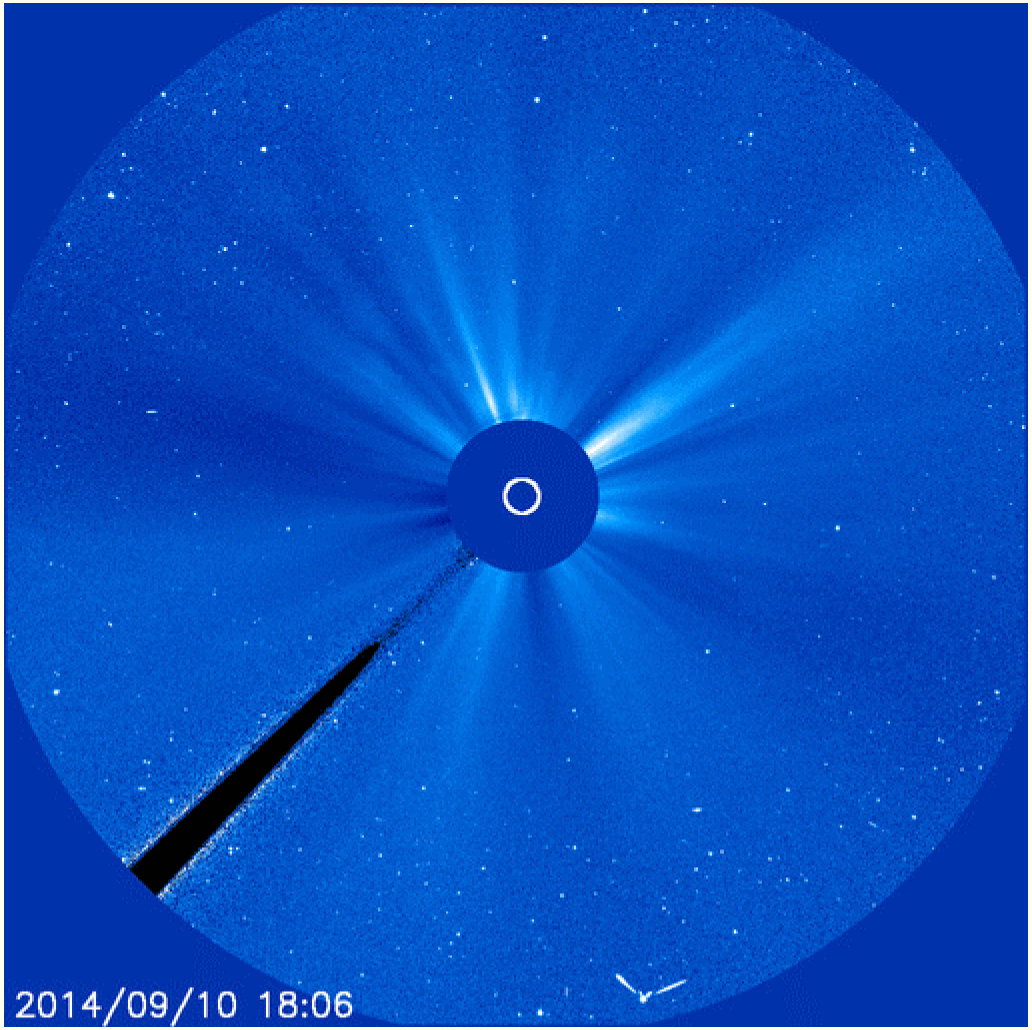
– **Det blir ingen** civilisationsomskakande händelse. Däremot kan det orsaka norrsken, säger han.

Enligt Dan Kiselman kan solstormar ge effekter på exempelvis kraftledningar men troligast är att vi bara märker av den här genom ett ljus på himlen. Norrsken är vanligt förekommande i de norra delarna av Sverige men i morgon och i övermorgon kan den som bor betydligt längre söderut se spännande fenomen på himlen nattetid.

Kan det synas ända ner mot Stockholmsområdet?

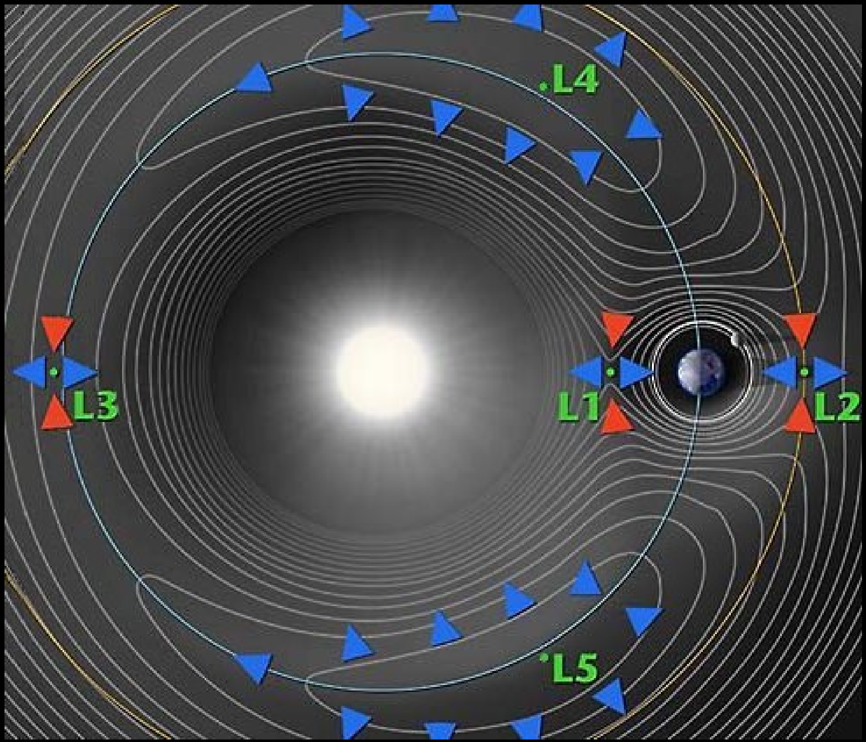
– Det är inte omöjligt. En extra titt på himlen skadar aldrig.

Coronal mass ejection, 2014-09-10



SOHO

(Solar and Heliospheric Observatory)

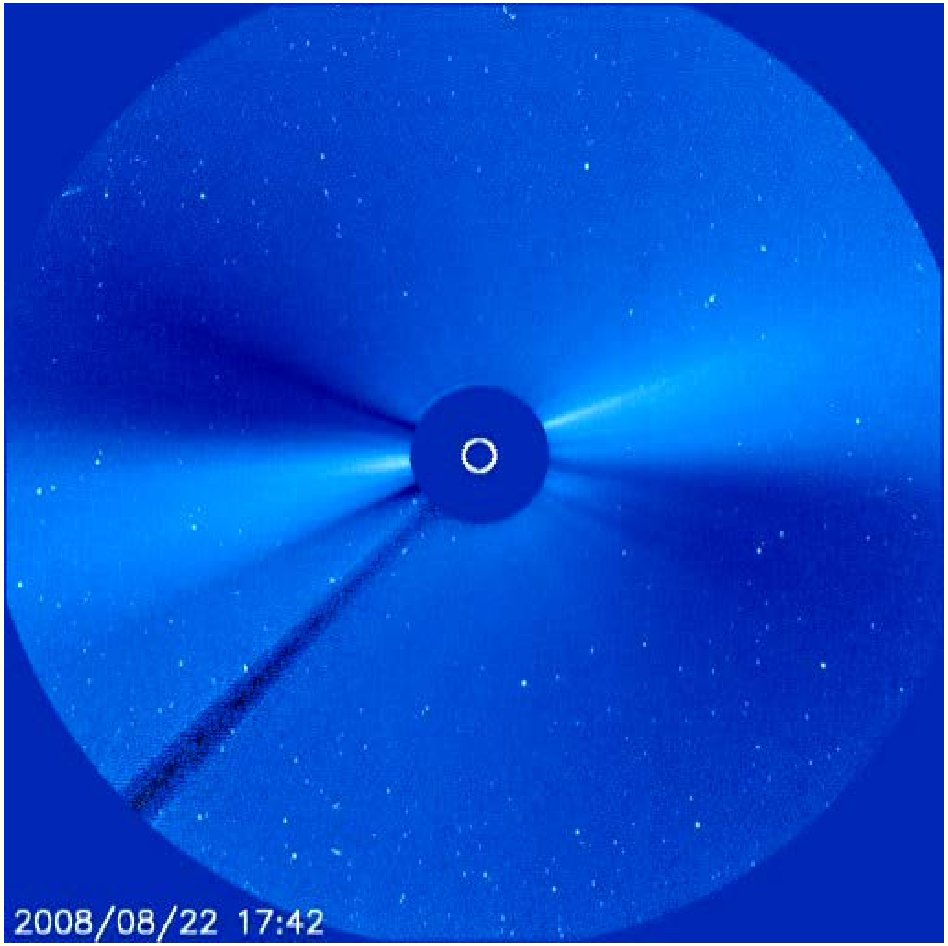


SOHO orbits the first Lagrange point

ESA - NASA collaboration

Solar wind

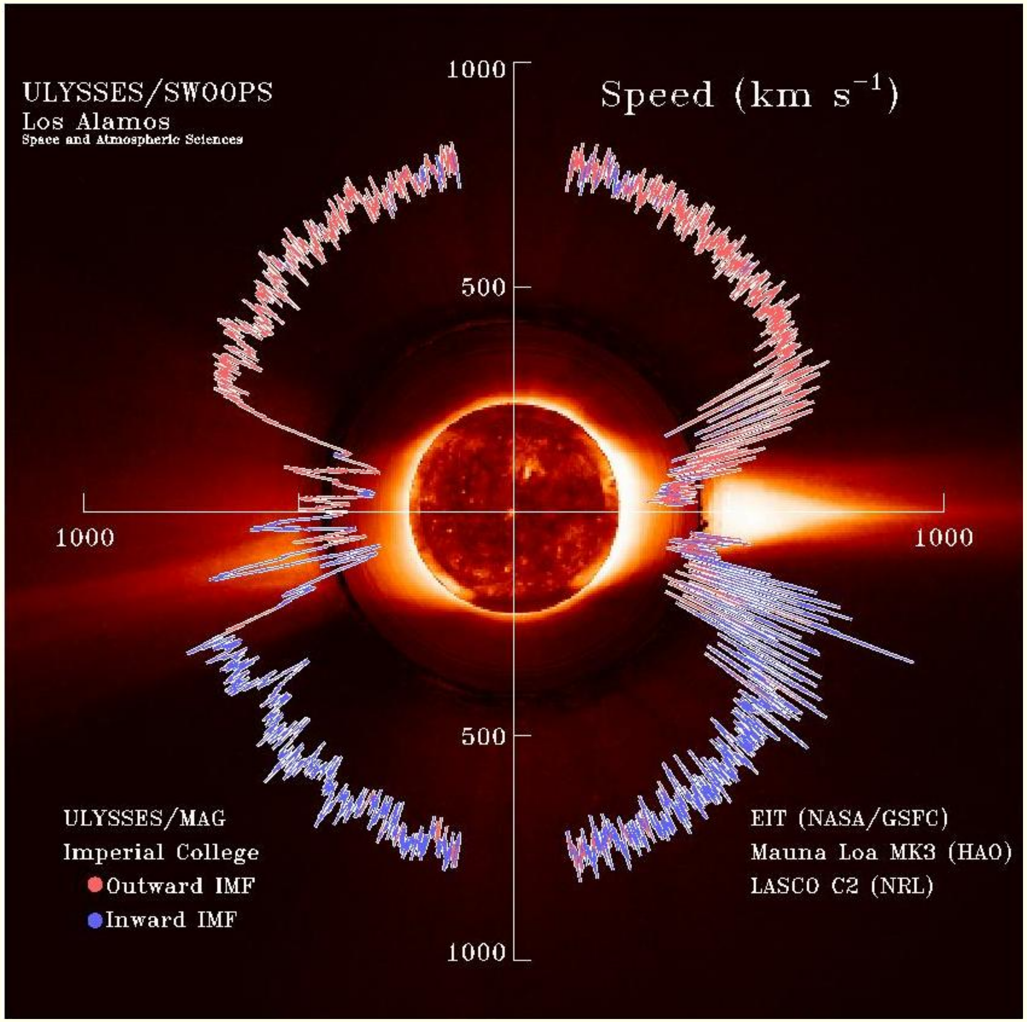
Corona continuously merges into solar wind



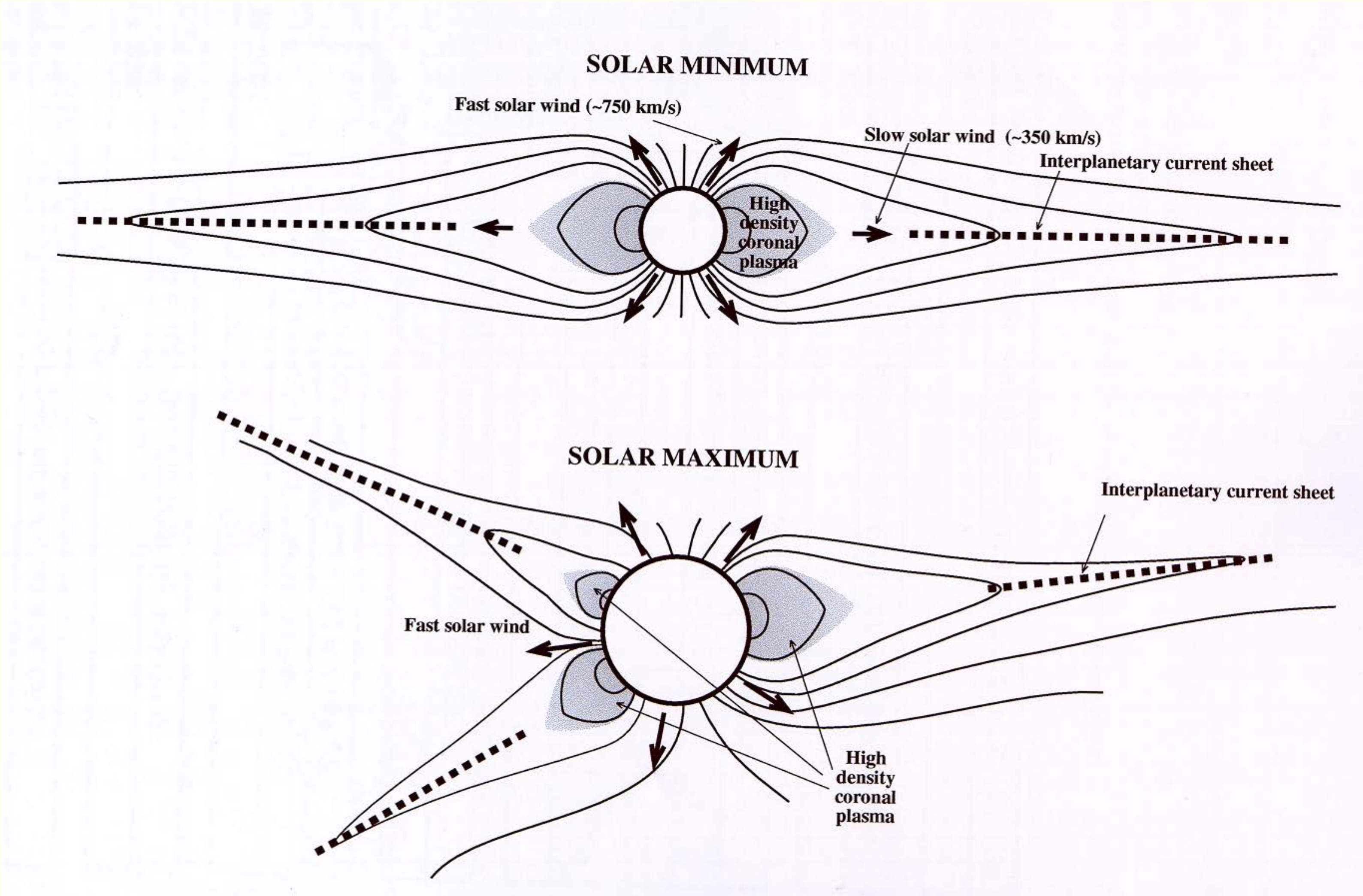
Solar and Heliospheric Observatory (SOHO)
LASCO C3 Coronagraph Movie

Solar wind

- **Fast solar wind in regions closer to poles**
- **Slow solar wind closer to equatorial plane**

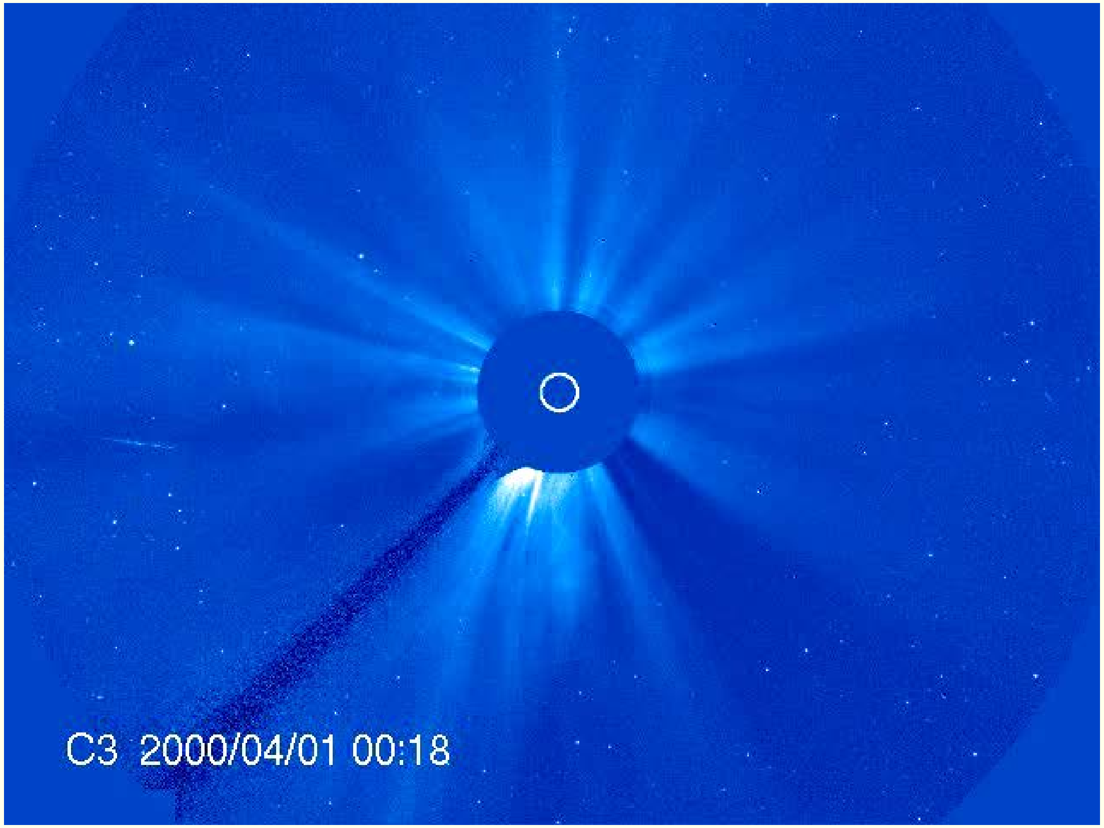


Solar wind



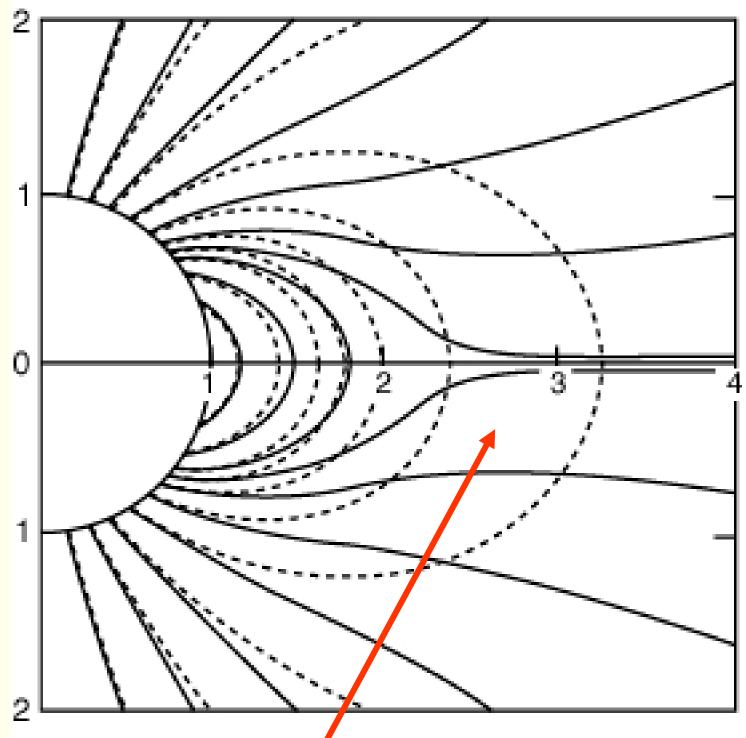
SOLAR MINIMUM
Fast solar wind (~750 km/s)
Slow solar wind (~350 km/s)
Interplanetary current sheet
High density coronal plasma
SOLAR MAXIMUM
Interplanetary current sheet
Fast solar wind
High density coronal plasma

More active solar wind



Solar and Heliospheric Observatory (SOHO)
LASCO C3 Coronagraph Movie

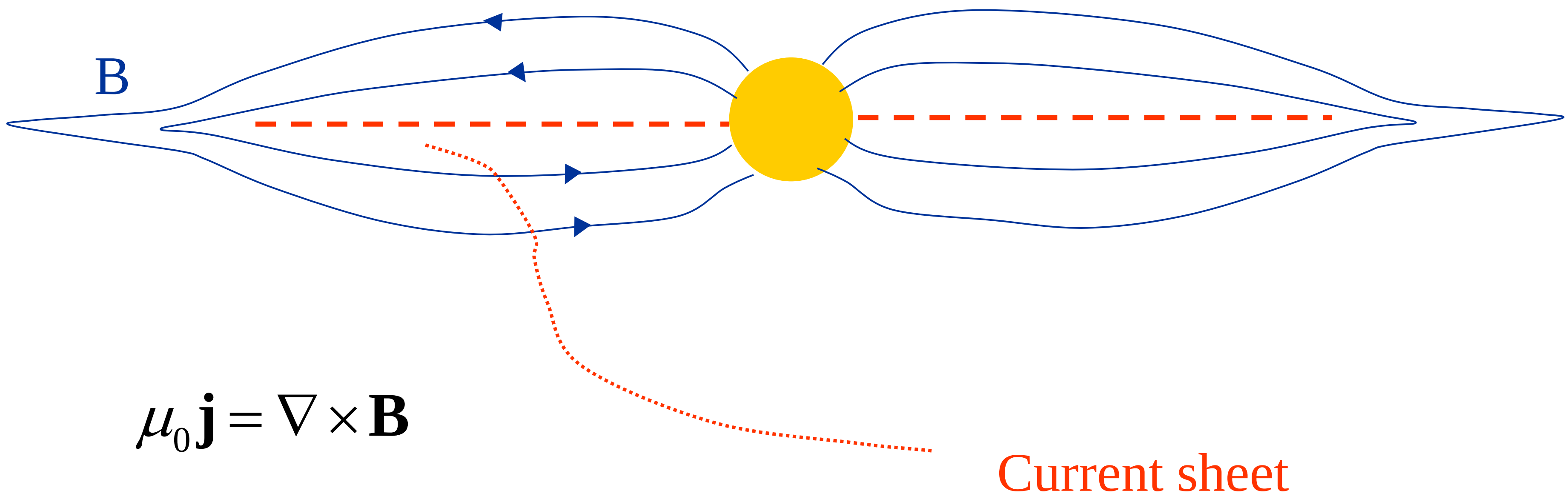
Helmet streamers



Magnetic field drawn out by solar wind.
This also brakes the solar wind.

Solar wind

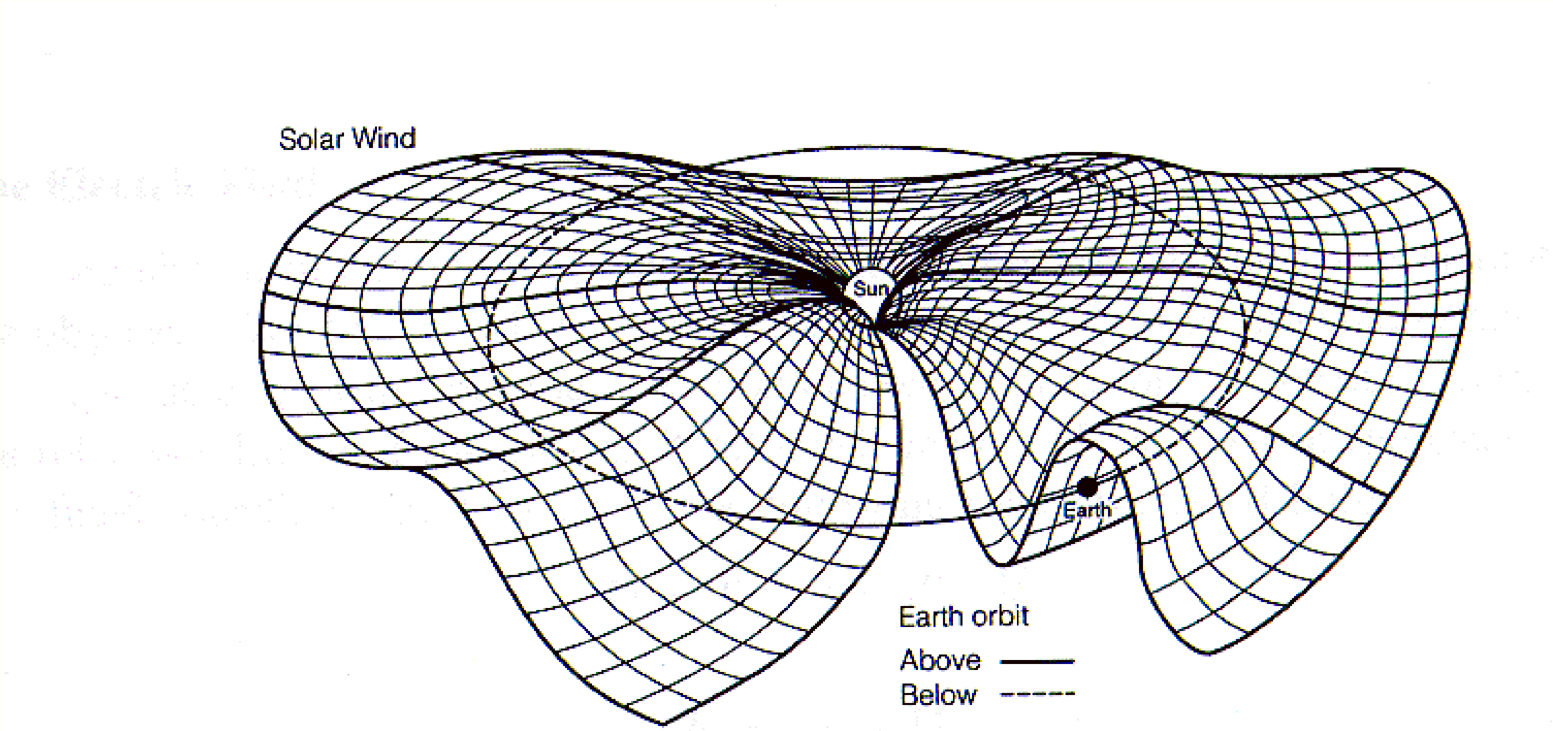
Interplanetary current sheet



$$\mu_0 \mathbf{j} = \nabla \times \mathbf{B}$$

Solar wind

Interplanetary current sheet



Later we will see that the N-S component of the interplanetary magnetic field (IMF is important for the coupling between solar wind and magnetosphere)

Solar wind

Some basic facts

Average values

$n_p = 8\ cm^{-3}$

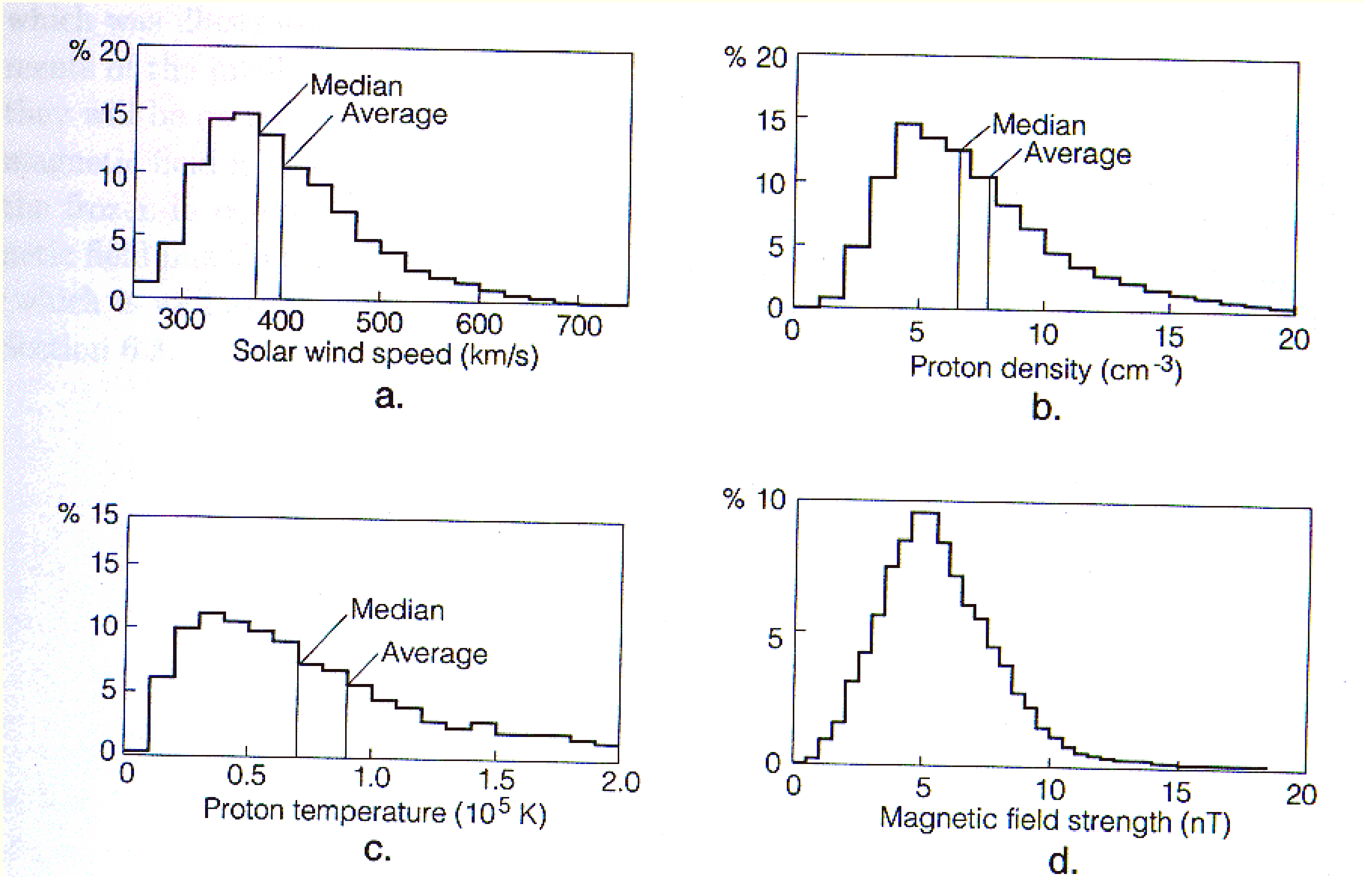
$v = 320\ km/s$

$T_p = 4 \cdot 10^4\ K$

$T_e = 10^5\ K$

$B = 5\ nT$

$\Phi_K = \rho v^3/2 = 0.22\ mW/m^2$



The solar wind today

Average values

$n_p = 8\ cm^{-3}$

$v = 320\ km/s$

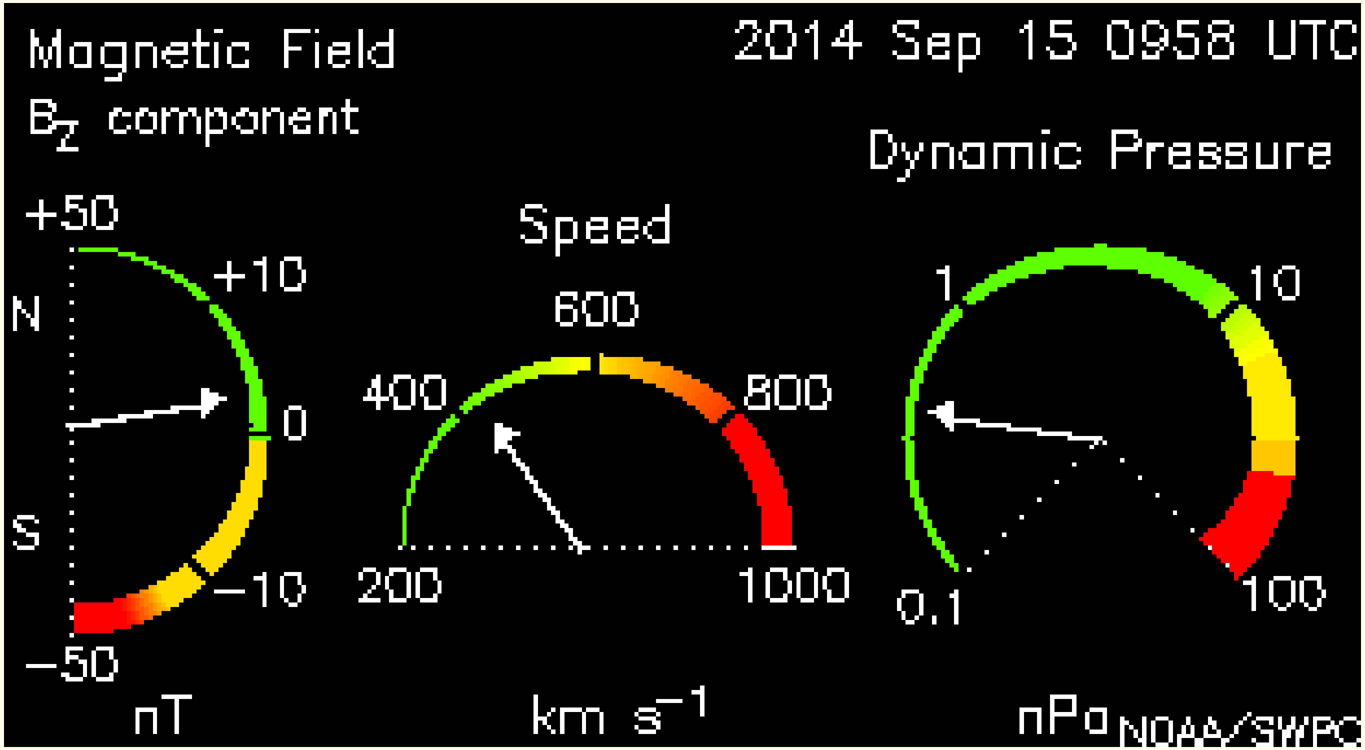
$T_p = 4 \cdot 10^4\ K$

$T_e = 10^5\ K$

$B = 5\ nT$

$p_D = \rho v^2/2 = 0.7\ nPa$

$\Phi_K = \rho v^3/2 = 0.22\ mW/m^2$



Measurements from ACE spacecraft

http://www.swpc.noaa.gov/SWN/

Space Weather Prediction Centre

Guess how long does it take the solar wind to flow from the Sun to the Earth?

Blue	8 min
Yellow	1.5 days
Green	5 hours
Red	5 days

$$t = \frac{s}{v} = \frac{1.496 \cdot 10^{11}}{320 \cdot 10^{3}} = 467\,500\ s = 129.9\ h = 5.4\ days$$

Red

But maybe

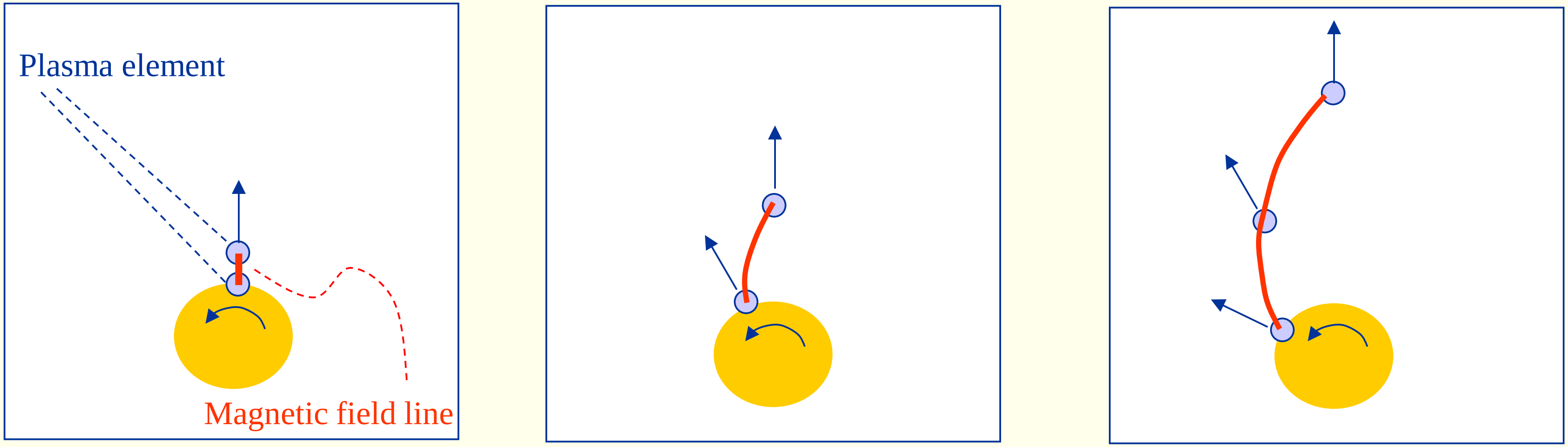
Yellow

if the solar wind is much faster

Does anyone happen to know the mathematical formula for the spiral caused by a rotating garden sprinkler?

Solar wind

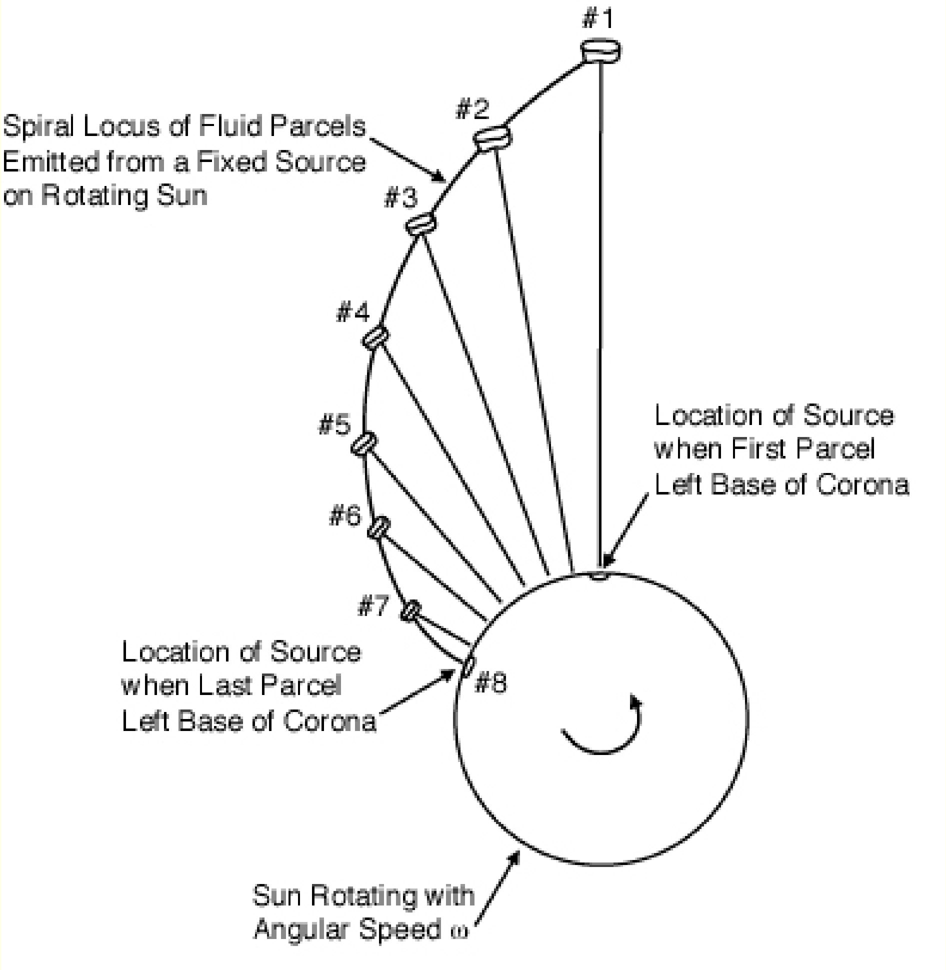
Magnetic field frozen into solar wind



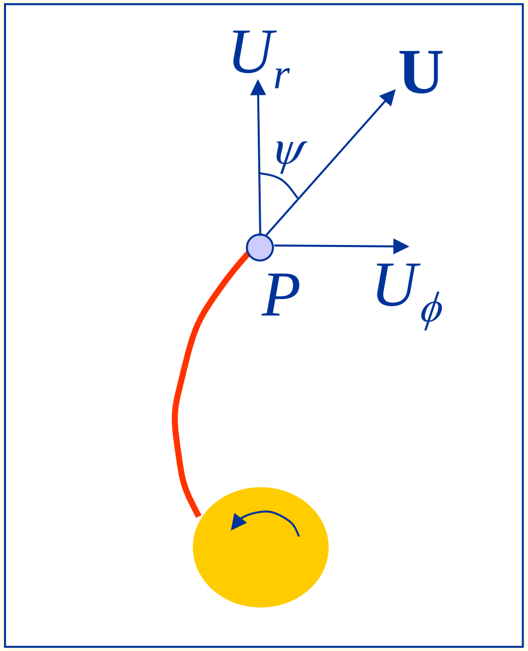
This is now seen from ”above”! (Looking down on the ecliptic plane from the pole.)

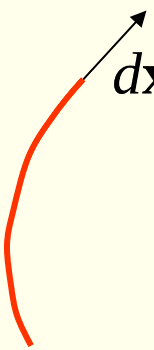
Solar wind

Parker spiral



Parker spiral





$$d\mathbf{x} = \mathbf{U}_{\mathbf{sw}} dt$$

Derivation of Ψ (Parker angle)

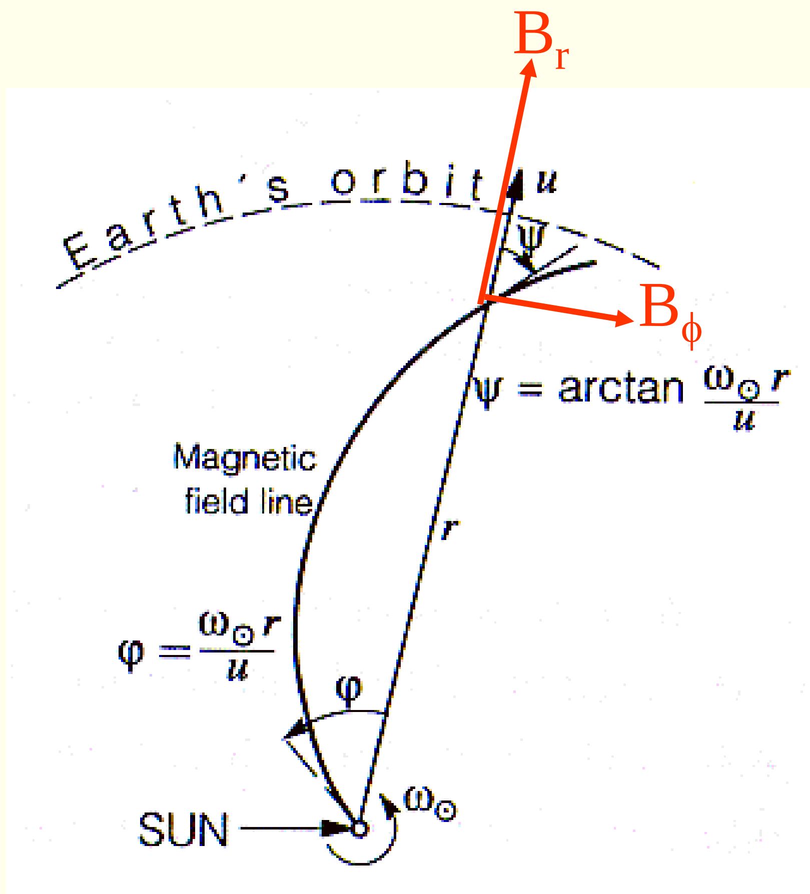
Consider a coordinate system rotating with the sun. The plasma element P in this coordinate system has two velocity components: U_r and U_ϕ .

Since the magnetic field is frozen into the solar wind, and follows the orbit of the plasma element P, at any time B has to be parallel to U. Then we have:

$$\tan\psi = \frac{B_\phi}{B_r} = \frac{U_\phi}{U_r} = \left(\frac{\omega r}{u_{SW}}\right)$$

Solar wind

Parker spiral



Archimedean spiral:

$$\frac{B_\phi}{B_r} = \tan\psi = \left(\frac{\omega r}{u_{SW}}\right)$$

Archimedean spiral

An Archimedean spiral (also arithmetic spiral), is a spiral named after the 3rd-century-BC Greek mathematician Archimedes; it is the locus of points corresponding to the locations over time of a point moving away from a fixed point with a constant speed along a line which rotates with constant angular velocity. Equivalently, in polar coordinates (r,ϕ) it can be described by the equation (*Wikipedia*)

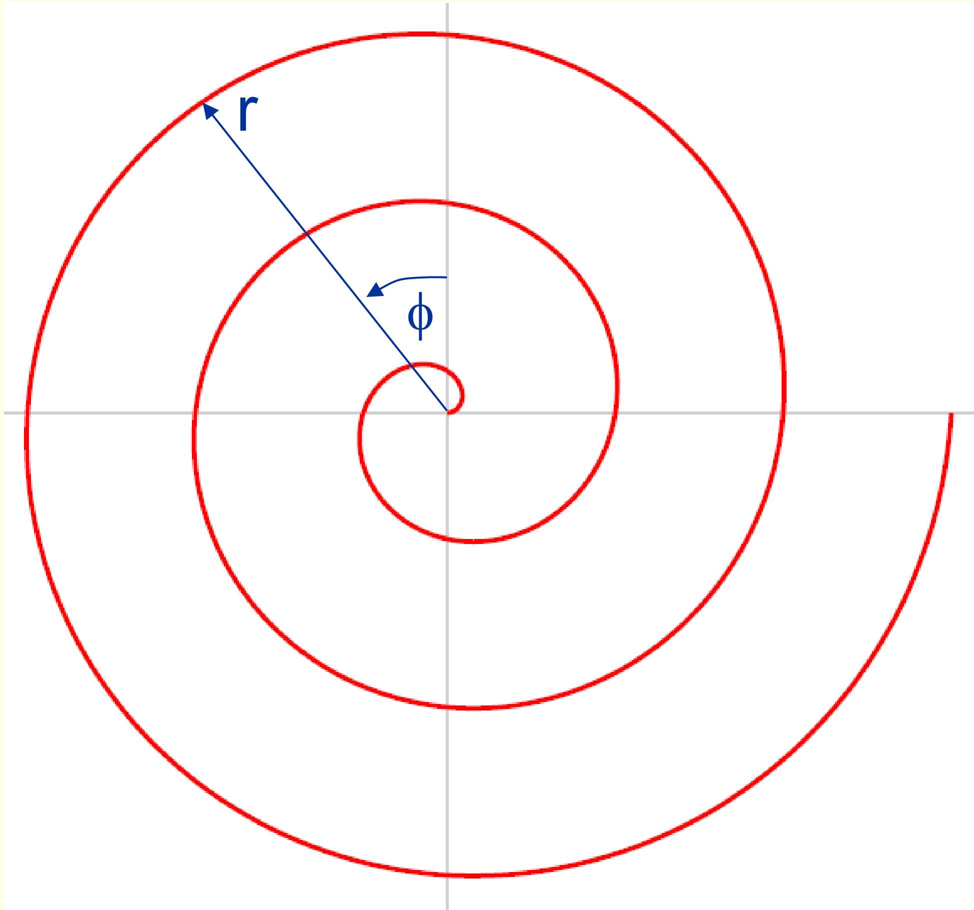
$$r = a + b\phi$$

$$r = a + b\omega t$$

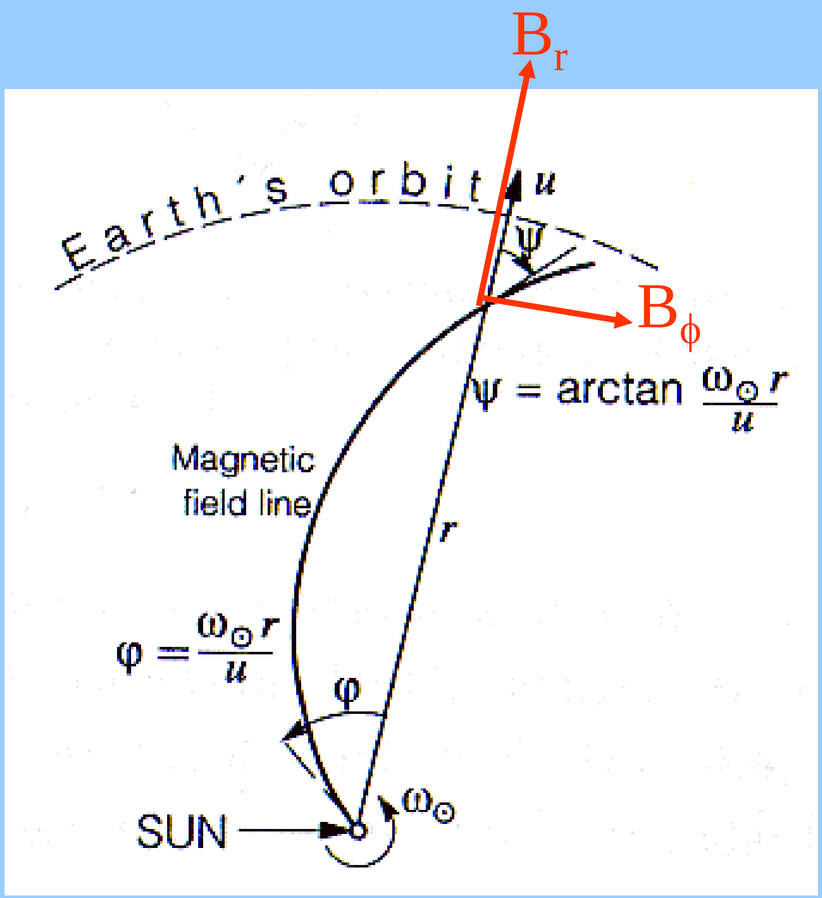
$$\frac{dr}{dt} = b\omega = u_{SW}$$

$$b = \frac{u_{SW}}{\omega}$$

$$r = R_{sun} + \frac{u_{SW}}{\omega}\phi$$



Use rotation period T of sun: $T = 27$ days



$r = 1$ A.U.

What is the angle Ψ at Earth's orbit for a typical solar wind speed?

Yellow	$\approx 50\,°$
Red	$\approx 80\,°$
Blue	$\approx 1\,°$
Green	$\approx 10\,°$

$$\Psi = \arctan(\frac{\omega r}{u})$$

What is ω? $$\omega = 2\pi f = \frac{2\pi}{T} = \frac{2\pi}{27 \cdot 24 \cdot 60 \cdot 60} = 2.7 \cdot 10^{-6} s^{-1}$$

$$\Psi = \arctan(\frac{\omega r}{u}) = \arctan(\frac{2.7 \cdot 10^{-6} \cdot 1.5 \cdot 10^{11}}{320 \cdot 10^{3}}) = \arctan(1.27) = 52^{\circ}$$

Yellow

Classification of plasmas

- **High density plasmas**
  - $\lambda << \rho$
  - *magnetic field not important, collisions dominate, isotropic.*

- **Medium density plasmas**
  - $\rho << \lambda << l_c$
  - *magnetic field important, collisions important, anisotropies.*

- **Low density plasmas**
  - $l_c << \lambda$
  - *magnetic field important, anisotropies, uninhibited motion along magnetic field*

ρ: *gyro radius*

λ: *mean free path*

l_c: *dimension of the plasma*

Plasma models/descriptions

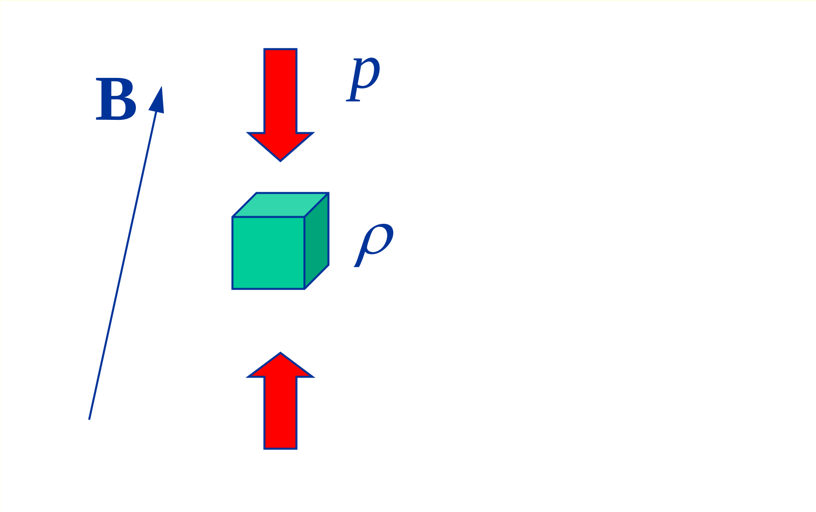
- Single particle motion
- Computer simulations of many-particle dynamics
- Generalization of statistical mechanics (kinetic theory)
- Generalization of fluid mechanics: *Magneto-hydrodynamics (MHD)*

Plasma physics
Magnetohydrodynamics (MHD)

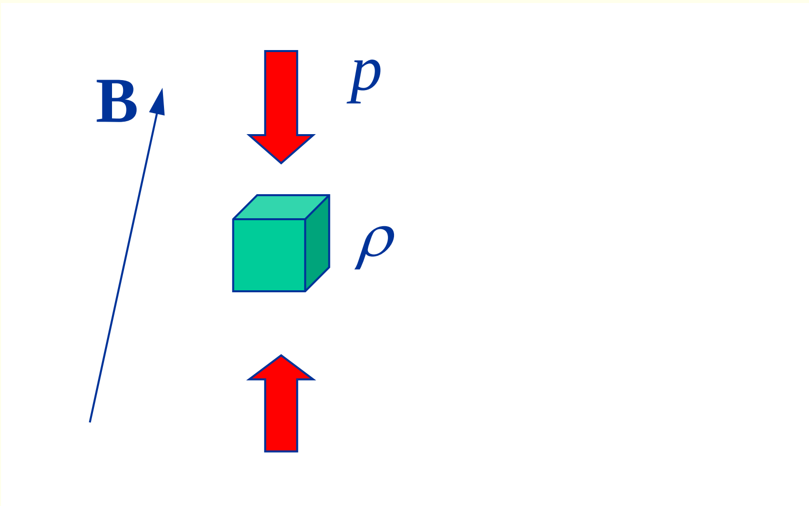
MHD is a combination of

- *fluid-/hydrodynamics* (which is based on Newton's laws of motion)
- *Maxwell's equations* (electrodynamics)

applied on a plasma volume element.



Magnetohydrodynamics (MHD)



For a volume element of plasma:

$$\mathbf{F} = m\mathbf{a} \quad \Rightarrow$$

quasineutrality

$$-\nabla p + n_e q \mathbf{v}_e \times \mathbf{B} + \rho_q \mathbf{E} = \rho \frac{d\mathbf{v}}{dt} \quad \Rightarrow$$

(1) $$\rho \frac{d\mathbf{v}}{dt} = -\nabla p + \mathbf{j} \times \mathbf{B}$$

Magnetohydrodynamics (MHD)

(1) $$\rho \frac{d\mathbf{v}}{dt} = -\nabla p + \mathbf{j} \times \mathbf{B}$$

This together with two of Maxwell's equations and Ohm's law make up the most common MHD equations:

(2) $$\mathbf{j} = \sigma(\mathbf{E} + \mathbf{v} \times \mathbf{B})$$

(3) $$\nabla \times \mathbf{B} = \mu_0 \left(\mathbf{j} + \varepsilon_0 \frac{\partial \mathbf{E}}{\partial t} \right)$$

Only consider slow variations

(4) $$\nabla \times \mathbf{E} = -\frac{\partial \mathbf{B}}{\partial t}$$

Magnetohydrodynamics (MHD)

(1) $$\rho \frac{d\mathbf{v}}{dt} = -\nabla p + \mathbf{j} \times \mathbf{B}$$

In equilibrium:

$$0 = -\nabla p + \mathbf{j} \times \mathbf{B} \iff$$

$$-\nabla p + \frac{1}{\mu_0}(\nabla \times \mathbf{B}) \times \mathbf{B} = 0$$

$$-\nabla p - \nabla\left(\frac{B^2}{2\mu_0}\right) + \frac{1}{\mu_0}(\mathbf{B} \cdot \nabla)\mathbf{B} = 0$$

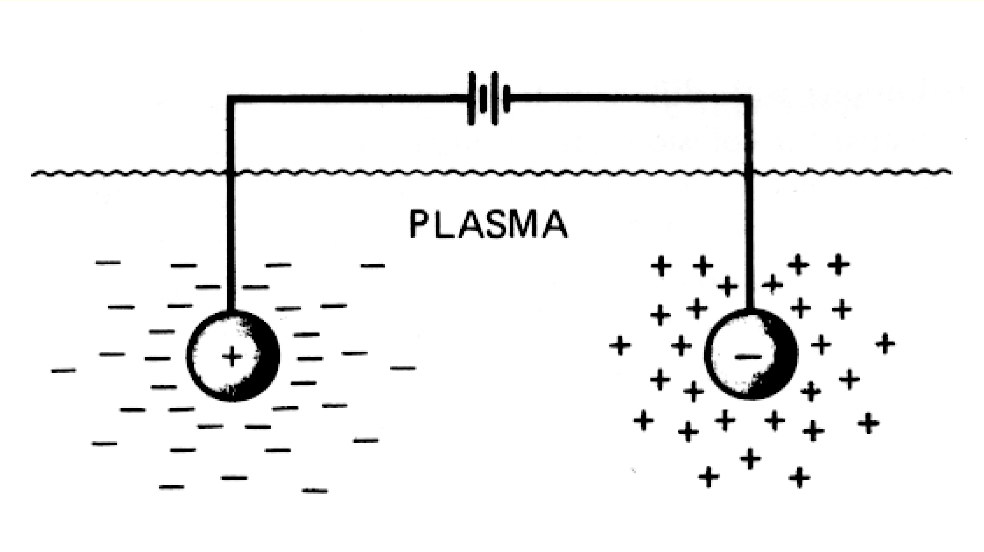
Represents tension in magnetic field

If magnetic tension = 0

$$p + \frac{B^2}{2\mu_0} = konst$$

Magnetic pressure

Quasineutrality



$$\Phi = \Phi_0 e^{-x/\lambda_D}$$

Debye length

$$\lambda_D = \sqrt{\frac{\varepsilon_0 k_B T}{n_e e^2}}$$

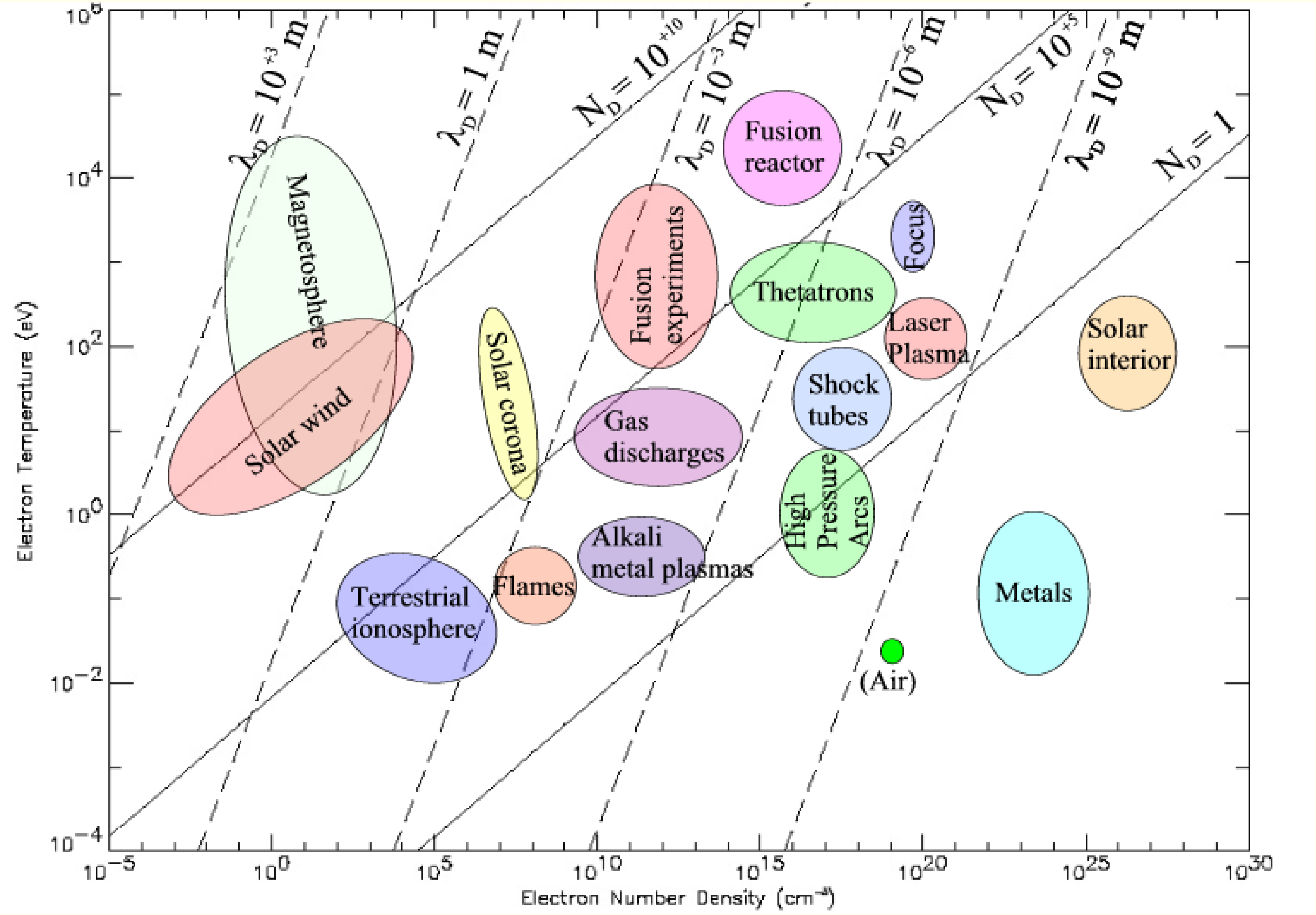
$$\frac{\Delta n}{n} = \frac{(n_e - n_i)}{n_e} < \left(\frac{\lambda_D}{l_c}\right)^2$$

$$l_C >> \lambda_D \quad \Rightarrow$$

Plasma close to neutral:

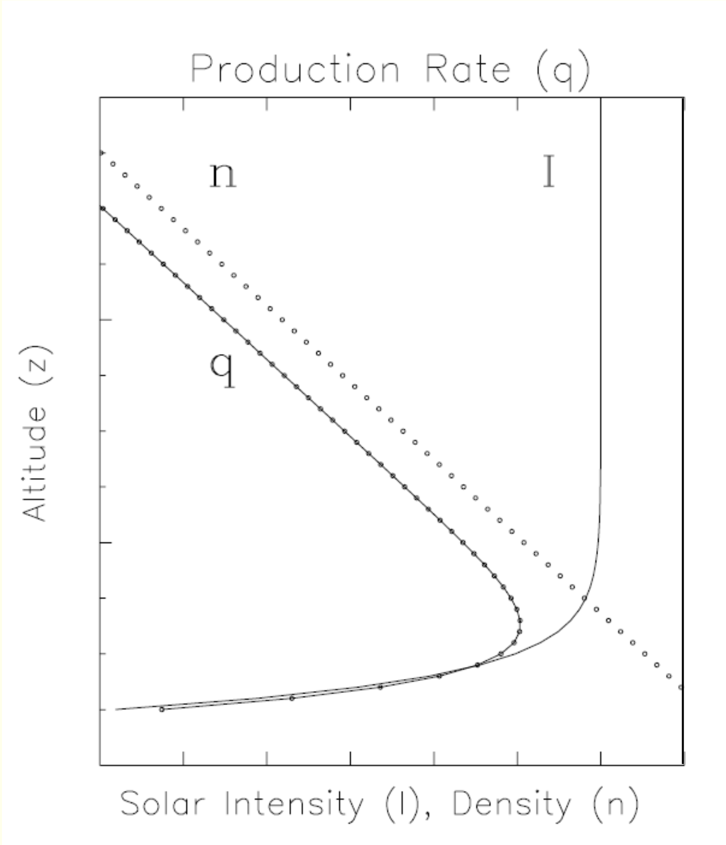
$$n_e \approx n_i$$

Debye lengths

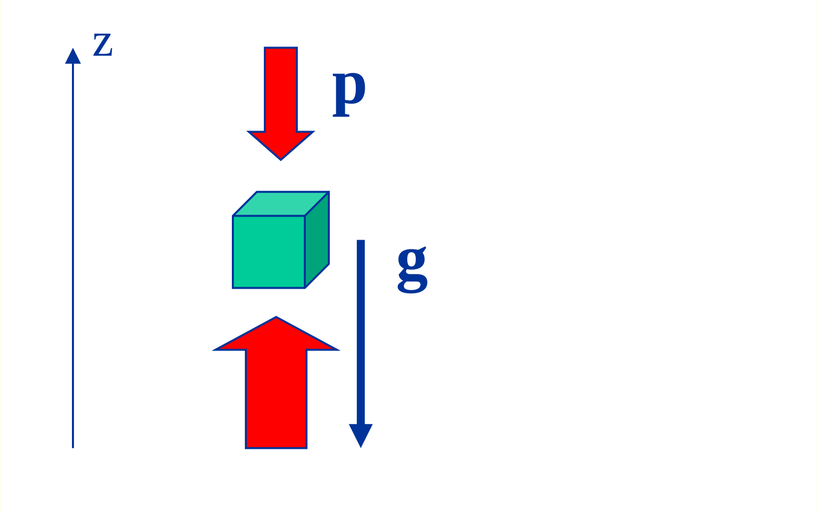


The ionosphere

Basic principle for creation of ionospheric layer



Atmospheric scale height



$$-\frac{dp}{dz} = g\rho_m$$ hydrostatic equilibrium for a volume element

$$p = nk_B T = \frac{\rho k_B T}{m}$$ ideal gas law

$$-\frac{k_B T}{m}\frac{d\rho_m}{dz} = g\rho_m$$ if T is constant

$$\rho_m = const \cdot e^{-z/(k_B T/gm)} = const \cdot e^{-z/H}$$

Scale height

$H = k_B T/gm$

Scale height

$$H = k_B T/gm$$

What is the approximate scale height in the atmosphere right here, right now?

(0° C = 273 K)

Blue	1 km	Yellow	30 km
Green	9 km	Red	100 km

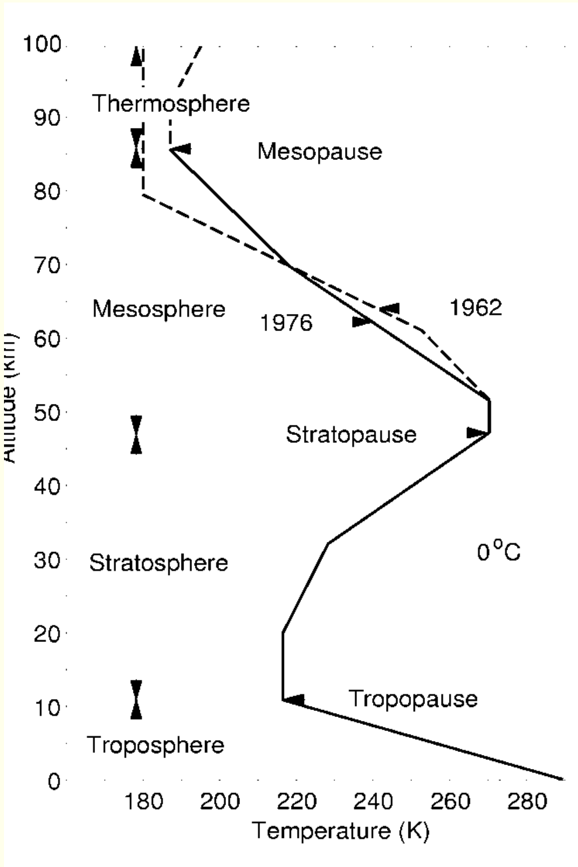
$$H = k_B T/gm = (1.38 \cdot 10^{-23} \cdot 290)/(9.81 \cdot 14 \cdot 2 \cdot 1.67 \cdot 10^{-27}) =$$

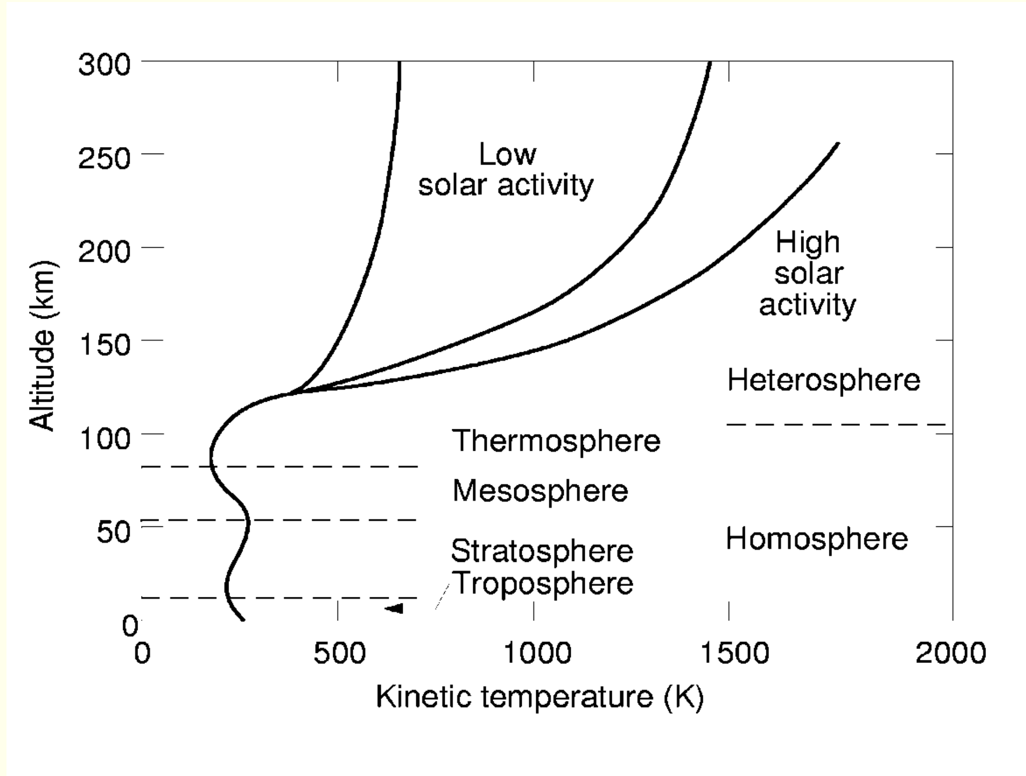
$$= 8724 \text{ m} \approx 9 \text{ km}$$

Green

What did we neglect when we derived the scale height?

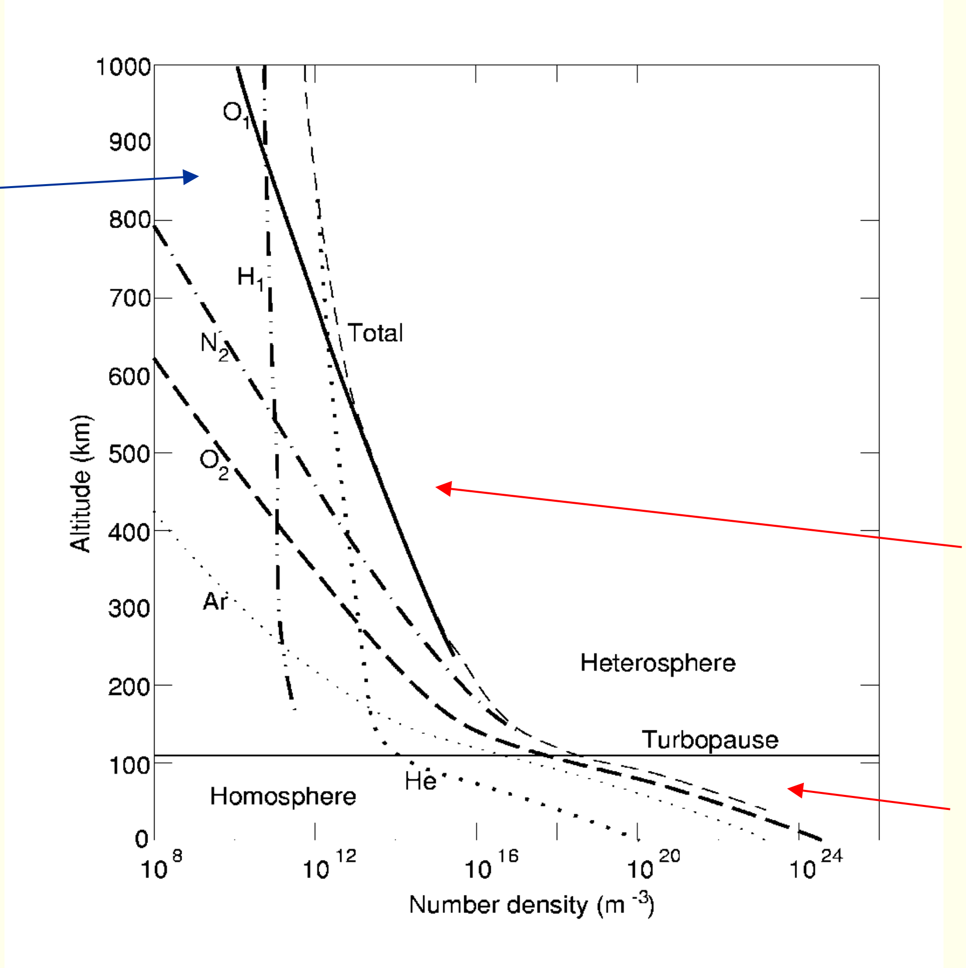
Temperature profile





Atmospheric composition

Longer scale height due to higher temperature



Separate scale heights for different components

Turbulent mixing – one scale height

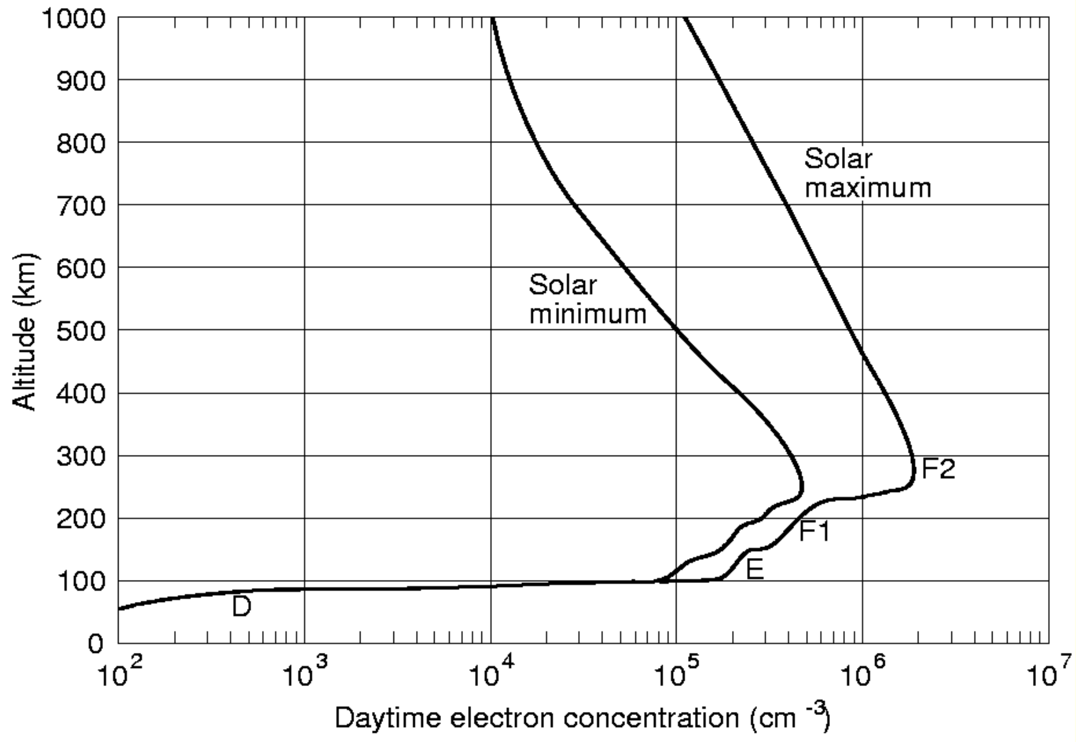
Ionosphere

- The ionized, electrically conducting part of the upper atmosphere
- The ionosphere is a **plasma**

History

- Stewart, 1882: Explained variations in the geomagnetic field
- Kenelly & Heavyside, 1902: explained Marconis transatlantic radio communication experiments
- Appleton & Barnett: experimental proof

Altitude distribution of electron density (n_e)



Continuity equation = conservation of ?

$$\frac{\partial n_e}{\partial t} = q - r - \nabla \cdot (n_e \mathbf{v}_e)$$

Ionization ($m^{-3}s^{-1}$)

Recombination ($m^{-3}s^{-1}$)

Flow ($m^{-3}s^{-1}$)

Continuity equation

$$\frac{dn_e}{dt} = q - r$$

$$q = a_i I n_n$$

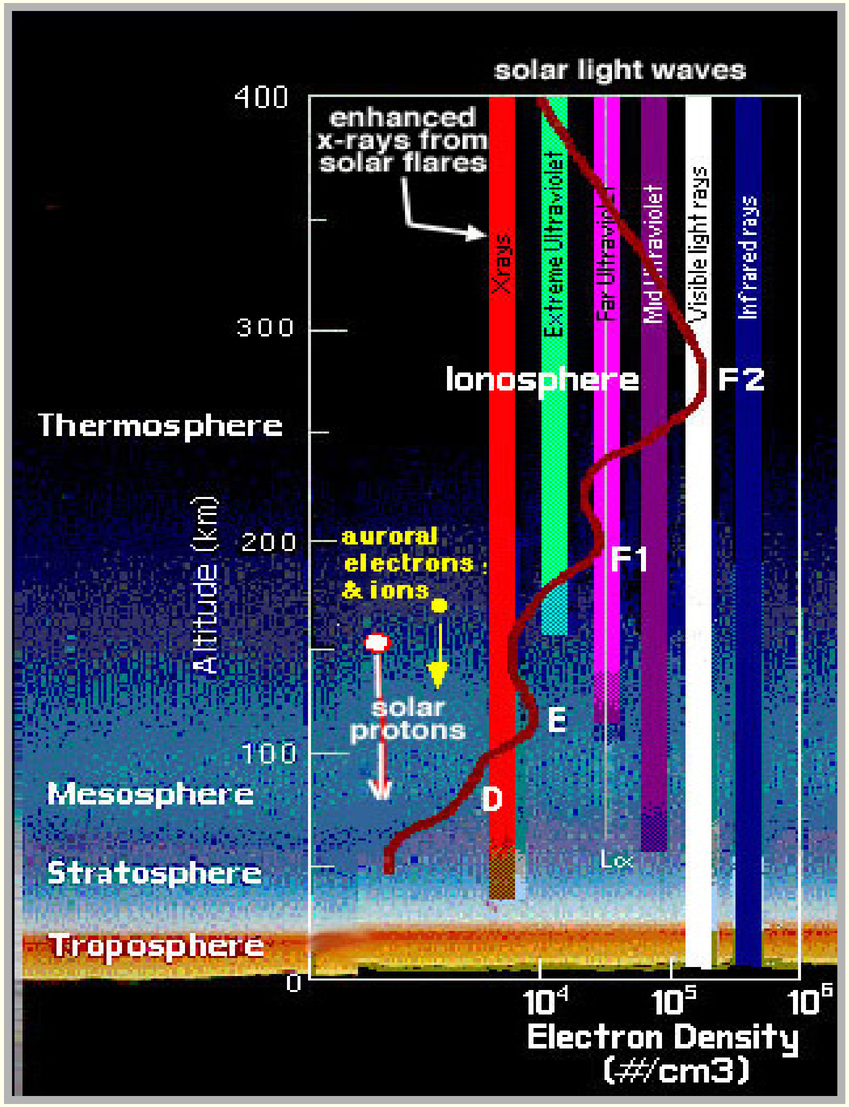
Ionization ($m^{-3}s^{-1}$)

Recombination ($m^{-3}s^{-1}$)

$$r = a_r n_e n_i = a_r n_e^2$$

Example: $e + O_2^+ \rightarrow O + O$ (dissociative recombination)

UV and X–ray radiation



$$\frac{dI}{dz} = I n_n a_a$$

Derive Chapman layer